

# Constraining the abundances of complex organics in the inner regions of solar-type protostars

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## ABSTRACT

The high abundances of Complex Organic Molecules (COMs) with respect to methanol, the most abundant COM, detected towards low-mass protostars, tend to be underpredicted by astrochemical models. This discrepancy might come from the large beam of the single-dish telescopes, encompassing several components of the studied protostar, commonly used to detect COMs. To address this issue, we have carried out multi-line observations of methanol and several COMs towards the two low-mass protostars NGC1333-IRAS2A and -IRAS4A with the Plateau de Bure interferometer at an angular resolution of 2 arcsec, resulting in the first multi-line detection of the O-bearing species glycolaldehyde and ethanol and of the N-bearing species ethyl cyanide towards low-mass protostars other than IRAS 16293. The high number of detected transitions from COMs (more than 40 methanol transitions for instance) allowed us to accurately derive the source size of their emission and the COMs column densities. The COMs abundances with respect to methanol derived towards IRAS2A and IRAS4A are slightly, but not substantially, lower than those derived from previous single-dish observations. The COMs abundance ratios do not vary significantly with the protostellar luminosity, over five orders of magnitude, implying that low-mass hot corinos are quite chemically rich as high-mass hot cores. Astrochemical models still underpredict the abundances of key COMs, such as methyl formate or di-methyl ether, suggesting that our understanding of their formation remains incomplete.

*Subject headings:* astrochemistry - ISM: abundances -ISM: individual objects (NGC 1333-IRAS2A, NGC 1333-IRAS4A) - ISM: molecules - stars: formation

## 1. Introduction

The early stages of low-mass star formation are known to be accompanied by the increase of the molecular complexity. Most of the lines detected in the sub-millimetric spectra of Class 0 protostars are attributed to Complex Organic Molecules (COMs, i.e. molecules based on carbon chemistry with 6 or more atoms; Herbst & van Dishoeck 2009), as shown by unbiased spectral surveys of low-mass protostars (see Caux et al. 2011, for the spectral survey of the low-mass protostar IRAS 16293-2422 for instance).

The bright protostars IRAS 16293-2422 and NGC1333-IRAS4A (hereafter IRAS 16293 and IRAS4A, respectively) have been the first two protostars where COMs, such as methyl formate, di-methyl ether, formic acid, methyl cyanide, or ethyl cyanide, have been detected with single-dish telescopes (Cazaux et al. 2003; Bottinelli et al. 2004a). The subsequent detection of a few COMs towards IRAS 16293 with interferometers, providing better angular resolutions of  $\sim 2''$ , by Bottinelli et al. (2004b) and Kuan et al. (2004) confirmed that most of the COMs emission likely comes from the warm inner region of protostellar envelopes, called “hot corinos”: the low-mass counterparts of high-mass hot cores. Since then, the number of low-mass protostars showing COMs has increased with the detection of COMs towards  $\sim 10$  other low-mass protostars by Bottinelli et al. (2007) and Öberg et al. (2011, 2014). Palau et al. (2011) and Fuente et al. (2014) also reported the detection of several COMs towards four low/intermediate-mass protostars.

Thanks to a larger number of transitions detected in the broad bands of their receivers, single-dish telescopes have been first used to derive the column densities of COMs, allowing them to constrain their abundances averaged over a relatively large beam of 10-30  $''$ .

The abundances of COMs, usually compared to that of their probable mother molecules (formaldehyde and methanol, see next paragraph), are found to be relatively high:  $\gtrsim 10$  % for methyl formate, di-methyl ether, and formic acid, and  $\sim 1$  % for methyl cyanide

and ethyl cyanide (see Bottinelli et al. 2007; Öberg et al. 2011, 2014) although the number of constrained abundance ratios remains relatively low and the abundances show some scattering between the sources and molecules.

Warm gas phase chemistry triggered by the sublimation of the main ice components, and of methanol in particular, has been first invoked by Millar et al. (1991) and Charnley et al. (1992) to explain the presence of COMs observed towards the high-mass hot core Orion KL by Blake et al. (1987). However, more recent laboratory experiments and theoretical calculations have contradicted several key assumptions made in the gas phase models: dissociative recombination of large ions do not lead predominantly to the formation of COMs but rather to their fragmentation into small pieces (Geppert et al. 2006; Hamberg et al. 2010) whilst ion-molecule reactions have been found to be not sufficiently efficient to produce the observed amount of methyl formate (Horn et al. 2004). The current scenario of COMs formation is now based on the recombination of radicals at the surface of interstellar grains during the warm-up phase ( $30\text{ K} < T < 100\text{ K}$ ) occurring in the envelopes surrounding Class 0 protostars (Garrod & Herbst 2006; Garrod et al. 2008). In these models, the radicals are generated by the UV photodissociation of the main ice components or they have survived to the incomplete hydrogenation process of CO leading to  $\text{CH}_3\text{OH}$  during the ice formation (Garrod & Herbst 2006; Taquet et al. 2012). Since the 90s, laboratory experiments have shown that UV irradiation of interstellar ice analogs containing methanol, formaldehyde, and ammonia can lead to the formation of a plethora of complex molecules and even amino-acids (Allamandola et al. 1988; Gerakines et al. 1996; Hudson & Moore 2000; Muñoz Caro et al. 2002; Öberg et al. 2009). However, the quantitative efficiency of the COMs formation in ices and their actual chemical pathways are still highly uncertain.

Although the current models produce a large set of COMs in significant quantities, they

are not able to explain the very high abundance ratio ( $> 10\%$ ) with respect to methanol seen for a few COMs, such as methyl formate or di-methyl ether (see Taquet et al. 2012, for a discussion of this problem). The discrepancy between observations and models could be due to the large beams of single-dish telescopes used to derive the abundance ratios of COMs. Typical single-dish beams of  $\sim 10''$  are much larger than the size of hot corinos ( $\sim 0.5''$ ; Maret et al. 2005; Maury et al. 2014) and encompass the cold envelopes and possible outflows driven by the central protostars, where COMs have also been detected (Arce et al. 2008; Öberg et al. 2010; Jaber et al. 2014). In addition, COMs have also been found in cold and quiescent cores by Bacmann et al. (2012) and Cernicharo et al. (2012) who claimed that quiescent cold gas phase chemistry can produce COMs but in lower quantities. However, new observations by Vastel et al. (2014) rather suggest that the emission from COMs observed in another pre-stellar core originates in an outer ring, so that the previous conclusions may need some cautions.

To better constrain the abundances of COMs originating from the hot corinos surrounding the low-mass protostars, a large number of transitions of methanol and COMs need to be observed with interferometers, providing angular resolutions of  $\sim 1\text{--}2''$ . The emission originating from hot corinos can be distinguished from other components of the envelope, allowing us to directly derive the abundance of COMs in the hot corinos. Moreover,  $\text{CH}_3\text{OH}$  emission is likely optically thick towards the continuum peak of protostars (Zapata et al. 2013), observations of its optically thin isotopologue  $^{13}\text{CH}_3\text{OH}$  are therefore required to derive an accurate estimate of the methanol column density. Although several publications have reported the interferometric detection of COMs in hot corinos (Bottinelli et al. 2004b; Kuan et al. 2004; Jørgensen et al. 2005; Bisschop et al. 2008; Jørgensen et al. 2011; Persson et al. 2012; Jørgensen et al. 2012; Maury et al. 2014), to our knowledge, none of them led to an accurate and simultaneous estimation of the column densities of methanol and COMs.

In this work, we present multi-line observations of methanol ( $^{12}\text{CH}_3\text{OH}$ , and  $^{13}\text{CH}_3\text{OH}$ ) as well as several COMs (methyl formate, di-methyl ether, ethanol, glycolaldehyde methyl cyanide, and ethyl cyanide) performed with the Plateau de Bure interferometer towards the two low-mass protostars IRAS2A and IRAS4A located in the NGC1333 star-forming region. Although the angular resolution does not allow us to spatially resolve the emission of COMs, it is sufficiently high to distinguish the emission from the hot corinos from other components of protostellar envelopes. The paper is structured as follows: section 2 describes the observational strategy, section 3 presents the continuum maps as well as the spectra and the maps of molecular transitions, section 4 explains the adopted methodology to derive the abundances of COMs, section 5 discusses the results; and section 6 summarizes this work with the conclusions.

## 2. Observations

The two low-mass Class 0 protostars IRAS2A and IRAS4A were observed with the IRAM Plateau de Bure Interferometer (PdBi) on 2010 July 20, July 21, August 1, August 3, November 24, and 2011 March 10 in the C and D configurations of the array. Due to the proximity to each other, the two sources were observed in the same track. Phase and amplitude were calibrated by performing regular observations of the nearby point sources 3C454.3, 3C84, and 0333+321. The amplitude calibration uncertainty is estimated to be  $\sim 20\%$ . The WIDEX backends have been used at 143.4 and 165.2 GHz, providing a bandwidth of 3.6 GHz each with a spectral resolution of 1.95 MHz ( $\sim 3.5 - 4 \text{ km s}^{-1}$ ). High-resolution narrow band backends focused on two  $\text{CH}_3\text{OH}$  lines and 12  $\text{HCOOCH}_3$  lines have also been used. They provide a bandwidth of 80 MHz with a spectral resolution of 0.04 MHz ( $0.08 \text{ km s}^{-1}$ ). Due to the low signal-to-noise ratio (S/N) obtained for the methyl formate lines at high spectral resolution, we decreased the spectral resolution to 0.4 MHz

(0.8 km s<sup>-1</sup>) to obtain a S/N ratio higher than 3. The data calibration and imaging were performed using the CLIC and MAPPING packages of the GILDAS software <sup>1</sup>. Continuum images were produced by averaging line-free channels in the WIDEX correlator before the Fourier transformation of the data. The coordinates of the source and the size of the synthesized beams are reported in Table 1.

### 3. Results

#### 3.1. Continuum Maps

Figure 1 shows the maps of the continuum emission of IRAS2A and IRAS4A at 143 and 165 GHz obtained after natural weighted cleaning. Parameters of the continuum emission (integrated flux and deconvolved FWHM size), obtained from elliptical Gaussian fits in the  $(u, v)$  plane, are given in Table 1. For the two settings, the FWHM size of the continuum emission is slightly smaller than the size of the synthesized beam, the continuum emission is consequently not resolved. In particular, IRAS4A is known to be a binary system with a 1.8'' separation (Looney et al. 2000), as depicted by the two red crosses in Fig. 1 that indicate the positions of IRAS4A-SE and -NW. Although the continuum emission of IRAS4A is peaked at the southeast (SE) position rather than at the northwest (NW) position for the two settings, we cannot resolve the two sources.

We estimated the mass of the envelope and the H<sub>2</sub> column density from the continuum fluxes and the sizes derived from elliptical Gaussian fits listed in Table 1. Assuming an optically thin dust emission, the mass  $M$  of the envelope is given by

$$M = \frac{S_\nu d^2}{K_\nu B_\nu(T_d) R_d} \quad (1)$$

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<sup>1</sup><http://www.iram.fr/IRAMFR/GILDAS>

Table 1: Properties of NGC 1333 IRAS2A and IRAS4A.

Source	IRAS2A	IRAS4A
R. A. (J2000)	03:28:55:57	SE: 03:29:10.52 NW: 03:29:10.42
Decl. (J2000)	31:14:37:22	SE: 31:13:31.06 NW: 31:13:32.04
d (pc) <sup>a</sup>	235	235
$V_{\text{LSR}}$ (km/s)	+7.7	+7.2
$L_{\text{bol}}$ ( $L_{\odot}$ ) <sup>b</sup>	36	9.1
$M_{\text{env}}$ ( $M_{\odot}$ ) <sup>c</sup>	5.1	5.6
Frequency = 143 GHz		
Beam size (")	2.2×1.7	2.1×1.7
Beam PA (°)	25	25
rms(WideX) <sup>d</sup>	2.57	3.34
rms(Cont.) <sup>e</sup>	1.56	10.8
Flux (Jy) <sup>f</sup>	0.13	1.1
Size (") <sup>f</sup>	1.7×1.7	2.1×1.7
PA (°) <sup>f</sup>	+51	+25
$M$ ( $M_{\odot}$ ) <sup>g</sup>	0.4	3.8
$N(\text{H}_2)$ ( $\text{cm}^{-2}$ ) <sup>g</sup>	$5.0 \times 10^{24}$	$3.7 \times 10^{25}$
Frequency = 165 GHz		
Beam size (")	2.3×1.7	2.4×1.8
Beam PA (°)	110	115
rms(WideX) <sup>d</sup>	3.50	4.02
rms(Cont.) <sup>e</sup>	1.84	10.8
Flux (Jy) <sup>f</sup>	0.19	1.6
Size (") <sup>f</sup>	1.9×1.7	2.5×1.6
PA (°) <sup>f</sup>	-67	-44
$M$ ( $M_{\odot}$ ) <sup>g</sup>	0.4	3.5
$N(\text{H}_2)$ ( $\text{cm}^{-2}$ ) <sup>g</sup>	$4.5 \times 10^{24}$	$3.1 \times 10^{25}$

Note. — *a*: Hirota et al. (2008); *b*: Karska et al. (2013); *c* : Kristensen et al. (2012); *d*: Units of mJy/beam/channel for a channel width of 1.95 MHz. *e*: Units of mJy/beam. *f*: Continuum integrated fluxes and sizes were obtained from elliptical Gaussian fits in the ( $u, v$ ) plane (i.e., deconvolved FWHM size); *g*: The envelope mass and averaged column density were derived from the continuum fluxes obtained within



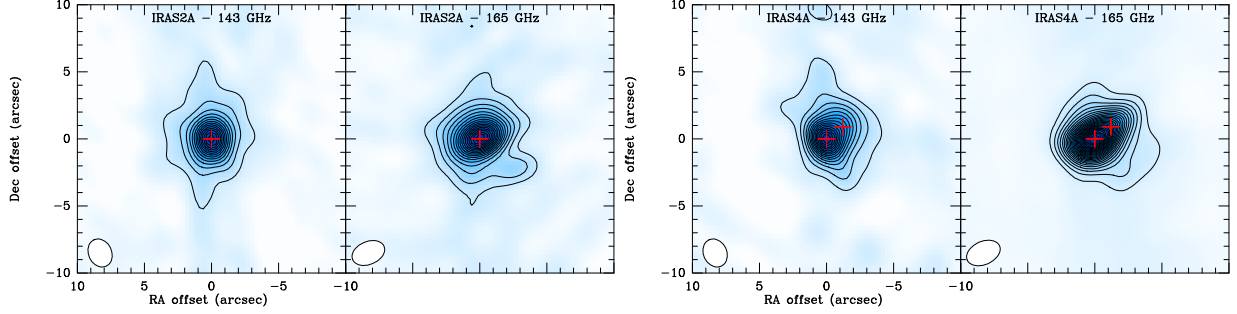


Fig. 1.— Continuum maps at 145 and 165 GHz of IRAS2A (rms of 1.56 and 1.84 mJy/beam, respectively) and IRAS4A (rms of 10.8 and 10.8 mJy/beam, respectively) . The positions of the sources IRAS2A, IRAS4A-NW, and IRAS4A-SE are marked by a red plus sign. Contour levels are in steps of  $3\sigma$ . The white ellipse represents the half-power beamwidth of the synthesized beam.

where  $S_\nu$  is the continuum flux integrated over the gaussian ellipse and listed in Table 1,  $d$  is the distance to the two low-mass protostars (235pc; Hirota et al. 2008),  $B_\nu(T_d)$  is the Planck black body function for a temperature  $T_d$  assumed to be 30 K,  $R_d$  is the dust-to-gas mass ratio equal to 0.01.  $K_\nu$  is the opacity per dust mass taken from column 5 in Table 5 of Ossenkopf & Henning (1994) (corresponding to grains showing a MRN size distribution covered by thin ice mantles at  $n_H = 10^6 \text{ cm}^{-3}$ ) and extrapolated to 145 GHz ( $K_\nu = 0.38 \text{ g/cm}^2$ ) and 165 GHz ( $K_\nu = 0.48 \text{ g/cm}^2$ ). The column density of  $\text{H}_2$  averaged over the gaussian ellipse can be deduced from  $M$  following this formula

$$N(\text{H}_2) = \frac{M}{\mu m_H \Omega d^2} \quad (2)$$

where  $M$  is the envelope mass,  $\mu = 2.38$  is the mean molecular mass in units of hydrogen atom masses,  $m_H$  is the hydrogen atom mass, and  $\Omega$  the solid angle subtended by the gaussian ellipse. The envelope mass  $M$  and  $\text{H}_2$  column density  $N(\text{H}_2)$  derived for the two frequencies are listed in Table 1.

### 3.2. Spectra

For the two sources and for all the settings, we obtained the spectral cubes by subtracting the continuum visibilities to the whole (line+continuum) datacube. For IRAS4A, the baseline has been flattened by importing the data cubes into CLASS and subtracting a polynomial function to each individual spectrum. The 3.6 GHz wide spectra obtained at the coordinates of IRAS2A and IRAS4A-NW with the two WideX backends are presented in Figure 2. We also present the narrow band spectra of  $\text{CH}_3\text{OH}$  and  $\text{HCOOCH}_3$  in Figures 3 and 4. The  $1\sigma$  rms noise in the line-free channels of all spectra are given in the caption of Fig. 2 and 3 and 4. The two sources are chemically rich since the WideX spectra towards IRAS2A and IRAS4A-NW display  $\sim 200$  and  $\sim 170$  lines detected with a signal-to-noise ratio higher than 3, resulting in a line density of 28 and 23 detected lines per GHz, respectively. For comparison, Maury et al. (2014) detected 86 lines above the 3-sigma level between 216.9 and 220.5 GHz towards IRAS2A with the PdBi, resulting in a similar line density of 24 detected lines per GHz.

The line identification was carried out using the JPL (Pickett et al. 1998) and the CDMS (Müller et al. 2005) spectroscopic catalogues. Line identifications were performed by eye, by taking into account the upper energy level, the line strength, and the velocity of each transition. The detected molecules are listed in Table 2 with the number of detected transitions, the energy range of their upper energy levels, and the spectroscopic reference. We detected about 35 transitions for the main isotopologue of methanol with upper energy levels up to  $\sim 1020$  K and about 13 lines for its isotopologue  $^{13}\text{CH}_3\text{OH}$ . In addition to methanol, we report the first multi-line detection of glycol aldehyde  $\text{HCOCH}_2\text{OH}$ , ethyl cyanide  $\text{C}_2\text{H}_5\text{CN}$ , and ethanol  $\text{C}_2\text{H}_5\text{OH}$  towards low-mass protostars other than IRAS 16293 <sup>2</sup>. We also detected several transitions originating from methyl formate  $\text{HCOOCH}_3$ ,

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<sup>2</sup> The detection of glycolaldehyde was reported in IRAS2A by Coutens et al. (2015) while

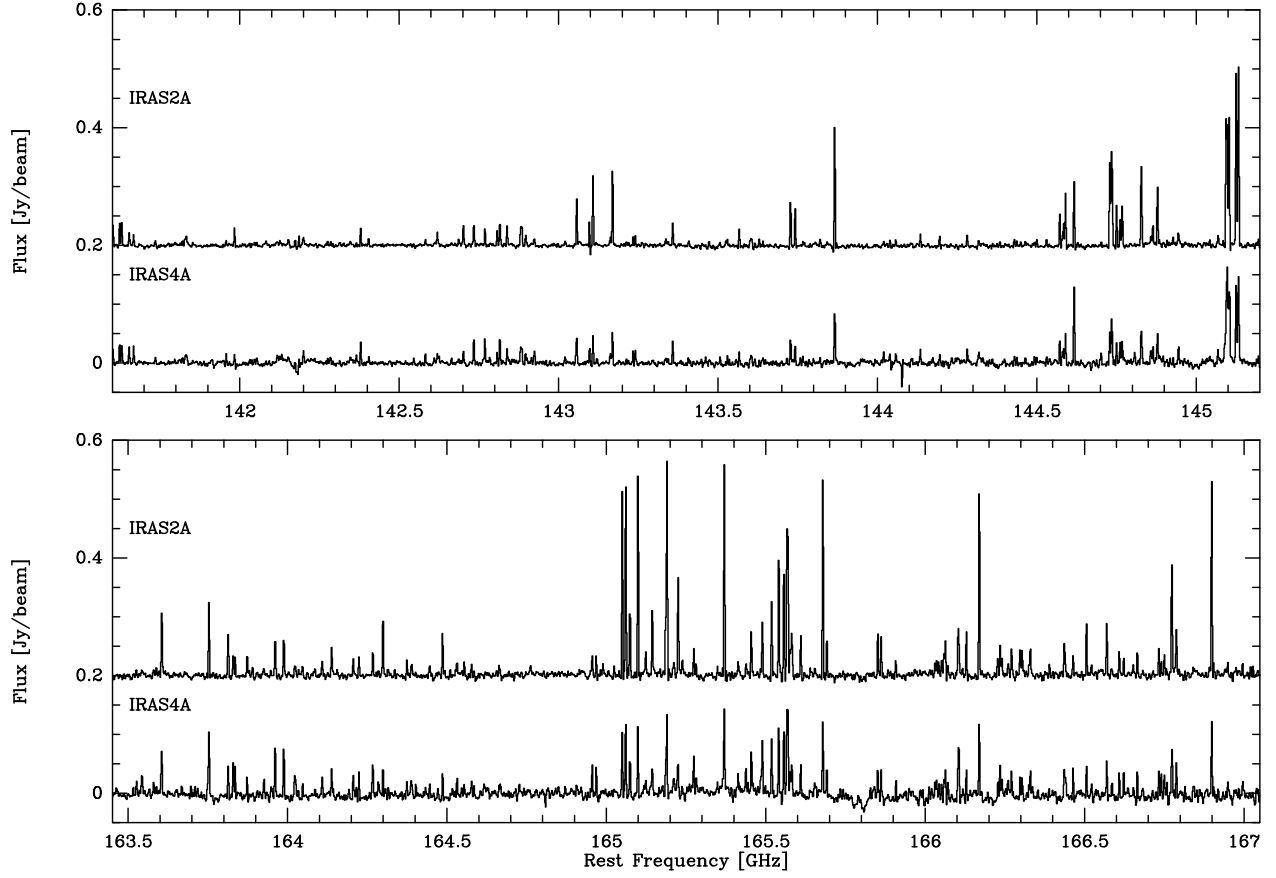


Fig. 2.— PdBi continuum-subtracted spectra of the WideX backends around 143 and 165 GHz towards the peak position of IRAS2A and IRAS4A-NW. The rms noise levels are 2.57 and 3.50 mJy/beam/channel at 143 GHz and 165 GHz towards IRAS2A and 3.34 and 4.02 mJy/beam/channel at 143.5 GHz and 165 GHz towards IRAS4A, respectively.

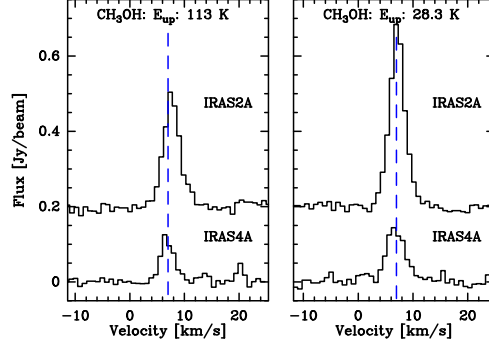


Fig. 3.— PdBi spectra of methanol obtained with the narrow band correlators towards the peak position of IRAS2A and IRAS4A-NW. The rms noise levels are 7.80 and 7.86 mJy/beam/channel towards IRAS2A and IRAS4A, respectively.

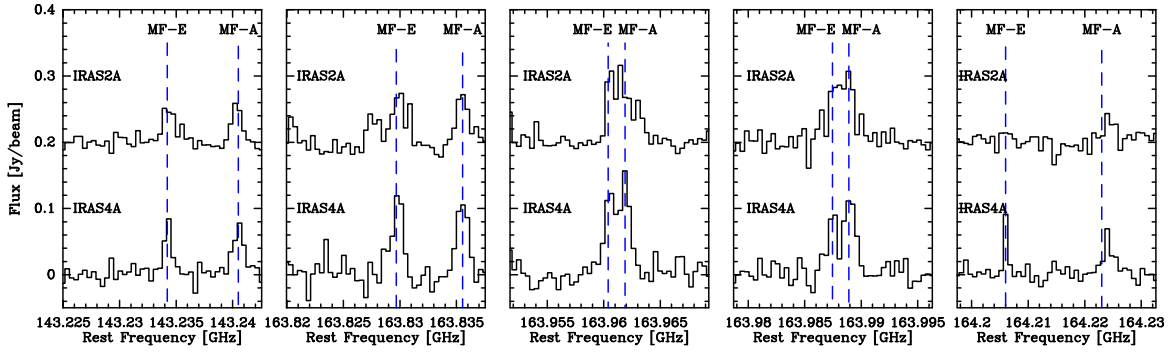


Fig. 4.— PdBi spectra of methyl formate obtained with the narrow band correlators towards the peak position of IRAS2A and IRAS4A-NW. The rms noises are about 7-10 mJy/beam/channel towards IRAS2A and IRAS4A.

di-methyl ether  $\text{CH}_3\text{OCH}_3$ , and methyl cyanide  $\text{CH}_3\text{CN}$  towards the two sources. Due to the low spectral resolution, we carefully checked that the detected lines do not suffer from any blending with other transitions from similar or other molecules by using the JPL and CDMS databases but also with the Splatalogue database <sup>3</sup>. In total, about 70 % of the lines detected at a 3 sigma level in the two 3.6 GHz WideX correlators are attributed to the complex organics listed in Table 2. Other lines are attributed to the deuterated methanol isotopologues  $\text{CH}_2\text{DOH}$ ,  $\text{CH}_3\text{OD}$ ,  $\text{CHD}_2\text{OH}$ , other deuterated molecules such as HDO (studied in a previous work; Taquet et al. 2013), DCN,  $\text{NH}_2\text{D}$ , and  $\text{DCO}^+$  (in absorption) and the sulphur-bearing species  $\text{SO}_2$ , and  $\text{C}^{34}\text{S}$ . The analysis of the deuterated methanol transitions will be published in a separate article.  $\sim 20$  lines are unidentified.

We estimated the Full-Width-at-Half-Maximum (FWHM) of the lines detected with the narrow band correlators giving a spectral resolution of 0.4 MHz, and with the WideX correlator providing a spectral resolution of 1.95 MHz, through a gaussian fit of the spectra obtained at the coordinates of IRAS2A and IRAS4A-NW. Table 3 lists the linewidths of the transitions detected with the two correlators. The uncertainties in the linewidths are due to the statistical errors from the gaussian fit and to the uncertainty from the low spectral resolution. In this work, the low spectral resolution dominates the uncertainty of the linewidths. Tables 8 to 16 of the Appendix list the properties of all detected transitions along with the FWHM linewidths. For the two sources, the FWHM of the  $\text{CH}_3\text{OH}$  transitions detected at high spectral resolution are about 3 km/s whilst the widths of the  $\text{HCOOCH}_3$  lines vary between 0.9 and 2.8 km/s when the lines of the -E and the -A states do not overlap. The linewidths derived with the PdBi are similar to the linewidths of other  $\text{CH}_3\text{OH}$  and  $\text{HCOOCH}_3$  transitions derived by Maret et al. (2005) and

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this article was in the review process.

<sup>3</sup><http://www.cv.nrao.edu/php/splat/>

Table 2: List of molecules detected towards IRAS2A and IRAS4A.

Molecule	IRAS2A		IRAS4A		Ref.
	$N_{\text{lines}}$	$E_{\text{up}}$	$N_{\text{lines}}$	$E_{\text{up}}$	
		(K)		(K)	
CH <sub>3</sub> OH	34	14 - 1022	35	14 - 1022	1
HCOOCH <sub>3</sub>	20	43 - 248	20	43 - 237	2
CH <sub>2</sub> DOH	13	33 - 230	13	33 - 230	3
<sup>13</sup> CH <sub>3</sub> OH	13	14 - 222	12	14 - 222	4
CH <sub>3</sub> OCH <sub>3</sub>	8	11 - 314	7	11 - 314	5
C <sub>2</sub> H <sub>5</sub> OH	8	37 - 216	7	37 - 216	6
CH <sub>3</sub> CN	7	40 - 390	7	40 - 390	7
CHD <sub>2</sub> OH	6	20 - 67	6	20 - 67	8
HCOCH <sub>2</sub> OH	4	53 - 68	7	53 - 177	9
SO <sub>2</sub>	4	24 - 102	4	24 - 102	10
C <sub>2</sub> H <sub>5</sub> CN	4	63 - 130	3	63 - 130	11
CH <sub>3</sub> OD	1	40	1	40	12
H <sub>2</sub> <sup>13</sup> CO	1	10	1	10	13
H <sub>2</sub> C <sup>18</sup> O	1	22	1	22	14
HC <sub>3</sub> N	1	75	1	75	15
NH <sub>2</sub> CHO	1	30	1	30	16
CH <sub>2</sub> CO	1	41	1	41	17
C <sup>34</sup> S	1	14	1	14	18
DC <sub>3</sub> N	1	62	-	-	15
HDO	1	319	1	319	19
DCN	1	10	1	10	20
NH <sub>2</sub> D	1	183	1	183	21
D <sub>2</sub> CO	1	21	1	21	14

Note. — 1: Xu et al. (2008); 2: Ilyushin et al. (2009); 3: Pearson et al. (2012); 4: Xu & Lovas (1997); 5: Lovas et al. (1979); 6: Endres et al. (2009); 7: Cazzoli and Puzzarini (2006); 8: Parise et al. (2002); 9: Carrol et al. (2010); 10: Alekseev et al. (1996); 11: Fukuyama et al. (1996); 12: Anderson et al. (1988); 13: Müller et al. (2000); 14: Dangoisse et al. (1978); 15: Lafferty & Lovas (1978); 16: Johnson et al. (1972); 17: Fabricant et al. (1977); 18: Bogey et al. (1982); 19: Messer et al. (1984); 20: De Lucia & Gordy (1969); 21: Cohen & Pickett (1982);

Bottinelli et al. (2004a). Due to the low spectral resolution, the linewidths derived from the WideX correlators result from the convolution of the intrinsic linewidths with the spectral resolution of 1.95 MHz. The top panel of Figure 5 shows the deconvolved FWHM of  $\text{CH}_3\text{OH}$ ,  $^{13}\text{CH}_3\text{OH}$ ,  $\text{HCOOCH}_3$ , and  $\text{CH}_3\text{CN}$ , showing a high number of transitions detected at a high signal-to-noise ratio, as a function of the energy of the upper level towards the two sources. For this purpose, we excluded several  $\text{CH}_3\text{OH}$  and  $\text{HCOOCH}_3$  transitions that are blended with each other. No clear trend can be deduced for the two sources. The fluctuation of the deconvolved FWHM linewidths between 2 and 8 km/s seems to be due to their high uncertainties. The bottom panel of Fig. 5 compares the linewidths deduced from the narrow band correlators giving a high spectral resolution of 0.4 MHz with the deconvolved linewidths from the WideX spectra for the lines detected with the narrow band correlators. For all the lines, the FWHM linewidth deconvolved from the WideX spectra is higher than the FWHM linewidth deduced from the narrow band correlator but the differences remain within the uncertainties.

### 3.3. Line Maps

For all the transitions, the interferometric maps of the IRAS2A and IRAS4A protostars have been obtained by integrating the flux over  $V_{\text{LSR}} \pm \Delta V$  where  $V_{\text{LSR}}$  is the system velocity of the source and  $\Delta V = 3$  km/s following the FWHM linewidths of the  $\text{CH}_3\text{OH}$  transitions listed in Table 3. In practice, due to the low resolution of the WideX backends, the line emission is integrated over 3 channels. Figures 6 and 7 show a compilation of the integrated line maps towards IRAS2A and IRAS4A obtained after natural weighted cleaning. For species where several transitions were detected, two maps showing a low-energy and a high-energy transition are presented. Tables 8 to 16 of the Appendix list the properties of all detected transitions along with their FWHM (full-width-at-half-maximum) sizes

Table 3: Linewidths of the transitions detected in the narrow band correlators.

Molecule	Frequency (GHz)	$E_{up}$ (K)	IRAS2A		IRAS4A	
			$dV_n$ (km/s)	$dV_W$ (km/s)	$dV_n$ (km/s)	$dV_W$ (km/s)
CH <sub>3</sub> OH	143.86580	28.3	$2.9 \pm 0.8$	$6.7 \pm 4.1$	$3.2 \pm 0.8$	$7.5 \pm 4.1$
CH <sub>3</sub> OH	143.16952	113	$3.2 \pm 0.8$	$6.1 \pm 4.1$	$2.8 \pm 0.8$	$5.8 \pm 4.1$
HCOOCH <sub>3</sub>	143.23420	47.3	$2.3 \pm 0.8$	$5.8 \pm 4.1$	$1.6 \pm 0.8$	$4.1 \pm 4.1$
HCOOCH <sub>3</sub>	143.24051	47.3	$2.8 \pm 0.8$	$5.8 \pm 4.1$	$1.9 \pm 0.8$	$4.1 \pm 4.1$
HCOOCH <sub>3</sub>	163.82968	62.5	$2.4 \pm 0.7$	$6.7 \pm 3.6$	$2.0 \pm 0.7$	$4.7 \pm 3.6$
HCOOCH <sub>3</sub>	163.83553	62.5	$1.9 \pm 0.7$	$3.6 \pm 3.6$	$2.4 \pm 0.7$	$4.4 \pm 3.6$
HCOOCH <sub>3</sub>	163.96039	64.5	$5.3 \pm 0.7$	$6.9 \pm 3.6$	$4.1 \pm 0.7$	$5.9 \pm 3.6$
HCOOCH <sub>3</sub>	163.96188	64.5				
HCOOCH <sub>3</sub>	163.98746	64.5	$5.2 \pm 0.7$	$6.9 \pm 3.6$	$0.9 \pm 0.7$	$6.0 \pm 3.6$
HCOOCH <sub>3</sub>	163.98891	64.5			$2.2 \pm 0.7$	
HCOOCH <sub>3</sub>	164.20598	64.9		$5.9 \pm 3.6$	$0.8 \pm 0.7$	$6.6 \pm 3.6$
HCOOCH <sub>3</sub>	164.22382	64.9	$2.1 \pm 0.7$	$6.5 \pm 3.6$	$2.3 \pm 0.7$	$5.2 \pm 3.6$

Note. —  $dV_n$  and  $dV_W$  are the non-deconvolved linewidths deduced from the spectra obtained with the narrow band and WideX correlators, respectively.



of emission and their Position Angle (PA) derived from the modelling of the visibilities assuming elliptical gaussians, circular gaussians or point sources if gaussian fits were not possible.

For the two sources, the emission of most molecules is limited to the inner regions near the protostars. In IRAS4A, the compact emission of all COMs transitions originates from the -NW source although the SE protostar is brighter in the continuum. In contrast, no molecular lines seem to originate from IRAS4A-SE, as also observed in previous interferometric observations of  $\text{H}_2^{18}\text{O}$  and other complex organics by Persson et al. (2012). The low-energy transitions of  $\text{CH}_3\text{OH}$  (the transition at  $\sim 145.094$  GHz including several non-resolved transitions towards IRAS2A not shown in this work and all the transitions with  $E_{up} \leq 120$  K towards IRAS4A) show an extended emission consistent with the position of molecular outflows. For IRAS2A, the interferometric map of the  $\text{CH}_3\text{OH}$  transitions at  $\sim 145.094$  GHz displays a slightly red-shifted emission located  $7''$  to the north from IRAS2A, consistent with the direction of an outflow previously detected at small scales with CO by Jørgensen et al. (2007). For IRAS4A, the emission of the low-energy  $\text{CH}_3\text{OH}$  lines extends in the bipolar outflow along a N-S direction and seems to peak at  $\sim 10''$  south from the protostars, consistent with the south lobe detected in SiO by Choi et al. (2005). From Fig. 7, it is clear that the inner hot corino does not dominate the flux of the weakly excited transitions of methanol in IRAS4A (see also Maret et al. 2005). The emission of the  $\text{HC}_3\text{N}$  transition is also spatially resolved and shows an elongation of its emission towards the SW direction for IRAS2A and towards the N direction for IRAS4A. The spatial distribution and the kinematics of the outflow driven by IRAS2A and IRAS4A will be analysed in a future work.

The emission originating from all COMs transitions but some  $\text{CH}_3\text{OH}$  and  $\text{HC}_3\text{N}$  lines is not spatially resolved by the array since their FWHM emission size deduced from the fit

to the visibilities is lower than the synthesized beam of the interferometer. Consequently, the source sizes presented in this work can only be used as upper limits. These observations are qualitatively consistent with previous models and observations suggesting that methanol and COMs mostly come from the inner hot corino. In this region, these molecules show a jump of their abundance when the temperature is higher than the temperature of ice sublimation ( $T \sim 100$  K). Maret et al. (2004) estimated a size for the hot corinos of IRAS2A and IRAS4A of 0.45 " by reproducing the formaldehyde emission observed with single-dish telescopes with an abundance jump of two orders of magnitude at  $r \sim 50$  AU from the central protostar. The luminosities assumed for IRAS2A and IRAS4A by Maret et al. (2004) are lower by a factor of 2.25 and 1.5, respectively, than the luminosities assumed in this work. Assuming that the temperature profile is governed by the Stefan-Boltzmanns law implies a difference in the temperature of a factor of 1.2 and 1.1, respectively. The assumption of higher luminosities for the two sources would therefore increase the size of the corino of a few AU. Maury et al. (2014) estimated a FWHM size of 0.4-0.9" for the hot corino of IRAS2A through the use of the more extended A configuration of the PdBi array.

## 4. Chemical Abundances in the Hot Corinos

### 4.1. Rotational Diagrams

First estimates of the excitation temperatures and the column densities of observed molecules in the hot corinos of IRAS2A and IRAS4A have been obtained from the rotational diagram (RD) analysis by assuming optically thin emission and a LTE population of the levels. Since the emission of most transitions is not spatially resolved with the PdBi, we measured the flux of all transitions originating from a circular mask with a diameter equal to the major axis of the synthesized beam size of the telescope ( $\sim 2.1$ - $2.6''$ ). Tables 8 to 16 of the Appendix list the measured flux for all transitions and in the two sources.

Astrochemical/dynamical models predict low abundances of COMs in the the dense regions of protostellar envelopes where the dust temperature is lower than the temperature of ice sublimation  $T_{\text{ev}}$  ( $\sim 100$  K) due to the efficient depletion, while the abundance profiles show a strong jump once  $T = T_{\text{ev}}$  (see Aikawa et al. 2008; Taquet et al. 2014) induced by thermal evaporation in the hot corino. We therefore assumed that all the flux measured in the  $\sim 2''$  mask comes from the hot corino region. Although the emission sizes of the COMs transitions are not necessarily similar (in IRAS 16293 for example, methyl formate has a source size of  $4''$  whereas formamide has a source size of  $2''$  Jaber et al. 2014) likely due to their different binding energies, we assumed a hot corino size  $\theta_s$  of  $0.5''$  for all the COMs observed towards the two sources (Maret et al. 2004). The linewidth at FWHM  $\Delta V$  is fixed to  $3$  km/s, which represents an average value of Table 3 and previous observations by Maret et al. (2005) and Bottinelli et al. (2007). For molecules with only a few lines detected within a narrow range of upper energy levels (glycolaldehyde, or ethyl cyanide) or for molecules showing only one detection, we assumed two values for the rotational temperature  $T_{\text{rot}}$ :  $T_{\text{rot}} = T_{\text{rot}}(\text{CH}_3\text{OH})$  and  $T_{\text{rot}} = T_{\text{evap}}(\text{ice}) = 100$  K. As already shown in several published observational works studying the emission of methanol towards high-mass and low-mass hot cores (Parise et al. 2006; Bisschop et al. 2007; Isokoski et al. 2013; Zapata et al. 2013), it is likely that low-energy transitions of  $\text{CH}_3\text{OH}$  are optically thick towards the center of protostellar envelopes, giving rise to an underestimation of their population in the rotational diagrams. Consequently, we excluded the  $\text{CH}_3\text{OH}$  transitions with  $E_{\text{up}} \leq 200$  K from our RD analysis. The rotational temperatures and total column densities of all species derived towards the two sources are summarized in Table 4.

For most molecules, observational data can be reasonably well fitted by a straight line with some scattering, likely due to opacity or non-LTE effects. For most of species, the column densities derived towards IRAS2A and IRAS4A are very similar. Column densities of  $\text{CH}_3\text{OH}$  are about  $6 - 12 \times 10^{17} \text{ cm}^{-2}$  whilst the column densities of complex organics

range between  $6 \times 10^{14}$  and  $6 \times 10^{16} \text{ cm}^{-2}$ . The rotational temperatures derived for most COMs are generally higher than 100 K in IRAS2A and in IRAS4A. Although they do not necessarily reflect the kinetic temperatures, the high rotational temperatures found in the inner protostellar envelope are in good agreement with the kinetic temperatures expected in hot corinos ( $T > 100 \text{ K}$ ). The  $\text{CH}_3\text{CN}$  population distribution can be used to estimate the kinetic temperature of the warm inner envelope because the  $\text{CH}_3\text{CN}$  population distribution over the different  $K$  ladders observed for  $J = 9$  can only be modified by collisions as radiative transitions are prohibited (see Wang et al. 2010, for a more detailed discussion of the  $\text{CH}_3\text{CN}$  population in hot cores). If the energy states are thermalized and the transitions are optically thin, the kinetic temperatures within the hot corinos of IRAS2A and IRAS4A would be close to 290 and 360 K, respectively. As we will see in the next section, the temperatures are probably overestimated since the  $\text{CH}_3\text{CN}$  transitions are likely optically thick.

We derived a  $^{12}\text{C}/^{13}\text{C}$  abundance ratio of  $\text{CH}_3\text{OH}$  of 26 and 12 in IRAS2A and IRAS4A, respectively. The two values are lower than the  $^{12}\text{C}/^{13}\text{C}$  abundance ratio of  $\sim 70$  expected in the local ISM (Boogert et al. 2002; Milam et al. 2005; Wiström et al. 2011) by a factor 2.7 and 5.8, respectively. We verified that using the same rotational temperature for  $\text{CH}_3\text{OH}$  and  $^{13}\text{CH}_3\text{OH}$  only modifies the  $^{12}\text{C}/^{13}\text{C}$  abundance ratios by a few percents at most. The low  $^{12}\text{C}/^{13}\text{C}$  abundance ratio might be due to the different ranges of excitation of the observed transitions used to derive the column densities. We used excited transitions (with  $E_{\text{up}} > 200 \text{ K}$ ) to derive  $N(^{12}\text{CH}_3\text{OH})$  whereas only weakly excited  $^{13}\text{CH}_3\text{OH}$  transitions (with  $E_{\text{up}} < 225 \text{ K}$ ) have been detected. The low ratio measured in the two protostars suggests that excited transitions of  $\text{CH}_3\text{OH}$  ( $E_{\text{up}} > 200 \text{ K}$ ) are also optically thick. An overestimation of the  $^{13}\text{CH}_3\text{OH}$  column density by the RD best fit, due to the large uncertainty on the fluxes, is also possible. An analysis taking the opacities into account is therefore required to clarify this issue.

Table 4: Results from the rotational diagram analysis.

Molecule	$N_{hc}^a$ ( $\text{cm}^{-2}$ )	$T_{\text{rot}}$ (K)	$X^b$	$X_{\text{meth}}^c$ (%)	$N_{hc}(\text{SD})^d$ ( $\text{cm}^{-2}$ )	$X_{\text{meth}}(\text{SD})^d$ (%)
IRAS2A						
CH <sub>3</sub> OH <sup>e</sup>	$(1.2 \pm 0.4) \times 10^{18}$	179 $\pm$ 62	$(2.5 \pm 0.9) \times 10^{-7}$		$1.4 \times 10^{17}$	
<sup>13</sup> CH <sub>3</sub> OH	$(4.8 \pm 1.3) \times 10^{16}$	164 $\pm$ 43	$(9.6 \pm 2.5) \times 10^{-9}$			
HCOOCH <sub>3</sub>	$(6.4 \pm 1.9) \times 10^{16}$	200 $\pm$ 61	$(1.3 \pm 0.4) \times 10^{-8}$	1.9 $\pm$ 0.8	$< 1.2 \times 10^{17}$	$< 85$
CH <sub>3</sub> CN	$(1.0 \pm 0.2) \times 10^{16}$	289 $\pm$ 63	$(2.0 \pm 0.4) \times 10^{-9}$	0.30 $\pm$ 0.10	$1.5 \times 10^{15}$	1
CH <sub>3</sub> OCH <sub>3</sub>	$(4.1 \pm 1.6) \times 10^{16}$	154 $\pm$ 62	$(8.2 \pm 3.3) \times 10^{-9}$	1.2 $\pm$ 0.6	$< 7.2 \times 10^{16}$	$< 53$
C <sub>2</sub> H <sub>5</sub> OH	$(5.1 \pm 2.2) \times 10^{16}$	325 $\pm$ 140	$(1.0 \pm 0.4) \times 10^{-8}$	1.5 $\pm$ 0.8	/	/
HCOCH <sub>2</sub> OH	7.8 $\times 10^{15}$	179 <sup>f</sup>	1.6 $\times 10^{-9}$	0.23 $\pm$ 0.06	/	/
	2.5 $\times 10^{15}$	100 <sup>f</sup>	5.0 $\times 10^{-10}$	0.074 $\pm$ 0.019		
C <sub>2</sub> H <sub>5</sub> CN	1.2 $\times 10^{15}$	179 <sup>f</sup>	2.4 $\times 10^{-10}$	0.036 $\pm$ 0.010	$< 1.7 \times 10^{16}$	$< 13$
	6.9 $\times 10^{14}$	100 <sup>f</sup>	1.4 $\times 10^{-10}$	0.021 $\pm$ 0.005		
HC <sub>3</sub> N	7.0 $\times 10^{14}$	179 <sup>f</sup>	1.4 $\times 10^{-10}$	0.021 $\pm$ 0.005	/	/
	7.1 $\times 10^{14}$	100 <sup>f</sup>	1.4 $\times 10^{-10}$	0.021 $\pm$ 0.006		
H <sub>2</sub> <sup>13</sup> CO	6.6 $\times 10^{15}$	179 <sup>f</sup>	1.3 $\times 10^{-9}$	0.20 $\pm$ 0.05	/	/
	2.1 $\times 10^{15}$	100 <sup>f</sup>	4.3 $\times 10^{-10}$	0.063 $\pm$ 0.017		
NH <sub>2</sub> CHO	1.2 $\times 10^{16}$	179 <sup>f</sup>	2.3 $\times 10^{-9}$	0.35 $\pm$ 0.09	/	/
	4.3 $\times 10^{15}$	100 <sup>f</sup>	8.7 $\times 10^{-10}$	0.13 $\pm$ 0.03		
CH <sub>2</sub> CO	7.0 $\times 10^{15}$	179 <sup>f</sup>	1.4 $\times 10^{-9}$	0.21 $\pm$ 0.05	/	/
	2.6 $\times 10^{15}$	100 <sup>f</sup>	5.2 $\times 10^{-10}$	0.077 $\pm$ 0.020		
IRAS4A						
CH <sub>3</sub> OH <sup>e</sup>	$(6.3 \pm 3.1) \times 10^{17}$	300 $\pm$ 151	$(1.7 \pm 0.9) \times 10^{-8}$		$2.0 \times 10^{17}$	
<sup>13</sup> CH <sub>3</sub> OH	$(5.1 \pm 1.5) \times 10^{16}$	197 $\pm$ 56	$(1.4 \pm 0.4) \times 10^{-9}$			
HCOOCH <sub>3</sub>	$(5.2 \pm 3.3) \times 10^{16}$	141 $\pm$ 90	$(1.4 \pm 0.9) \times 10^{-9}$	1.5 $\pm$ 1.0	$1.0 \times 10^{17}$	52
CH <sub>3</sub> CN	$(6.5 \pm 2.9) \times 10^{15}$	360 $\pm$ 162	$(1.8 \pm 0.8) \times 10^{-10}$	0.18 $\pm$ 0.10	$2.6 \times 10^{15}$	1
CH <sub>3</sub> OCH <sub>3</sub>	$(3.1 \pm 1.0) \times 10^{16}$	86 $\pm$ 27	$(8.5 \pm 2.6) \times 10^{-10}$	0.87 $\pm$ 0.37	$< 4.5 \times 10^{16}$	$< 22$
C <sub>2</sub> H <sub>5</sub> OH	$(4.4 \pm 1.4) \times 10^{16}$	221 $\pm$ 69	$(1.2 \pm 3.7) \times 10^{-9}$	1.2 $\pm$ 0.5	/	/
HCOCH <sub>2</sub> OH	$(8.9 \pm 3.4) \times 10^{15}$	124 $\pm$ 48	$(2.4 \pm 0.9) \times 10^{-10}$	0.25 $\pm$ 0.12	/	/
C <sub>2</sub> H <sub>5</sub> CN	2.3 $\times 10^{15}$	300 <sup>f</sup>	6.2 $\times 10^{-11}$	0.064 $\pm$ 0.018	$1.9 \times 10^{15}$	$< 0.92$
	8.2 $\times 10^{14}$	100 <sup>f</sup>	2.2 $\times 10^{-11}$	0.023 $\pm$ 0.007		
HC <sub>3</sub> N	6.7 $\times 10^{14}$	300 <sup>f</sup>	1.8 $\times 10^{-11}$	0.019 $\pm$ 0.005	/	/
	6.8 $\times 10^{14}$	100 <sup>f</sup>	1.8 $\times 10^{-11}$	0.019 $\pm$ 0.005		
H <sub>2</sub> <sup>13</sup> CO	5.3 $\times 10^{15}$	300 <sup>f</sup>	1.4 $\times 10^{-10}$	0.15 $\pm$ 0.04	/	/
	1.1 $\times 10^{15}$	100 <sup>f</sup>	3.1 $\times 10^{-11}$	0.032 $\pm$ 0.009		
NH <sub>2</sub> CHO	8.5 $\times 10^{15}$	300 <sup>f</sup>	2.3 $\times 10^{-10}$	0.24 $\pm$ 0.07	/	/
	2.1 $\times 10^{15}$	100 <sup>f</sup>	5.7 $\times 10^{-11}$	0.059 $\pm$ 0.017		
CH <sub>2</sub> CO	1.6 $\times 10^{16}$	300 <sup>f</sup>	4.2 $\times 10^{-10}$	0.43 $\pm$ 0.12	/	/
	4.0 $\times 10^{15}$	100 <sup>f</sup>	1.1 $\times 10^{-10}$	0.11 $\pm$ 0.03		

Note. —

a: Column densities averaged over a source size of 0.5 ''(see text). b: The abundances relative to H<sub>2</sub> are obtained from  $N(\text{H}_2)$  derived at 145 GHz in Table 1 assuming an homogeneous H<sub>2</sub> column density within the beam. c: The abundance ratios are relative to the <sup>13</sup>CH<sub>3</sub>OH column density multiplied by 70 (see text). d: Column densities derived from previous single-dish observations and scaled to a source size of 0.5 ''. CH<sub>3</sub>OH: Maret et al. (2005), other COMs in IRAS2A: Bottinelli et al. (2007), other COMs in IRAS4A: Bottinelli et al. (2004a). e:  $N_{hc}$  and  $T_{\text{rot}}$  have been derived from the rotation diagram neglecting the transitions with  $E_{\text{up}} < 105$  K (see text). f: When they could not be derived from the rotation diagram, the rotational temperatures have been assumed to be equal to the temperature of evaporation of water ice and the rotational temperature of CH<sub>3</sub>OH.

## 4.2. Population Diagrams

We used the so-called Population Diagram (PD) analysis following the method described by Goldsmith & Langer (1999) to investigate the effect of optical depth on the column densities of each level. We applied the PD analysis where four or more transitions were detected for each species. Briefly, the PD analysis includes the influence of optical depths on the level populations assumed to be at LTE, following the formula

$$\ln\left(\frac{N_{\text{up}}}{g_{\text{up}}}\right) = \ln\left(\frac{N_{\text{tot}}}{Q_{\text{rot}}}\right) - \frac{E_{\text{up}}}{kT_{\text{rot}}} - \ln\left(\frac{\Omega_a}{\Omega_s}\right) - \ln(C_\tau) \quad (3)$$

where  $N_{\text{tot}}$  is the total column density of the species in question,  $N_{\text{up}}$  is the observed column density of the upper state of the species with an upper energy  $E_{\text{up}}$  including the opacity effect,  $Q_{\text{rot}}$  is the partition function,  $\Omega_a$  is the beam solid angle,  $\Omega_s$  is the source solid angle.  $C_\tau$  is given by

$$C_\tau = \frac{\tau}{1 - \exp(-\tau)} \quad (4)$$

with  $\tau$  being the optical depth.  $\tau$  can be expressed as

$$\tau = \frac{c^3}{8\pi\nu_0^3} \frac{A_{\text{ul}}}{\Delta V} \frac{g_{\text{up}} N_{\text{tot}}}{Q_{\text{rot}}} \exp\left(\frac{-E_{\text{up}}}{kT_{\text{rot}}}\right) \left(\exp\left(\frac{h\nu_0}{kT_{\text{rot}}}\right) - 1\right) \quad (5)$$

where  $c$  is the speed of light,  $A_{\text{ul}}$  is the Einstein-A coefficient of spontaneous emission,  $\Delta V$  is the FWHM fixed to 3 km/s. We performed a reduced  $\chi_{\text{red}}^2$  minimization by running a grid of 125,000 models covering a large parameter space in rotational temperature  $T_{\text{rot}}$  (50 values between 10 and 500 K), total column density in the source  $N_{\text{tot}}$  (50 values between  $10^{15}$  and  $10^{20}$  cm $^{-2}$ ), and source size  $\theta_s$  (50 values between 0.04 and 2 "). At LTE, the column density of every upper state  $N_{\text{up}}$  can be derived for each set of  $N_{\text{tot}}$ ,  $T_{\text{rot}}$ , and source solid angle  $\Omega_s = \theta_s^2$  according to equation (3). The best-fit model populations are plotted together with the observed populations of the levels in Figures 9 and 10 and are marked by red cross symbols. Tables 5 and 6 summarise the parameters of the best-fit models.

We started the PD analysis by simultaneously modelling the population distribution of  $^{12}\text{CH}_3\text{OH}$  and  $^{13}\text{CH}_3\text{OH}$ . For this purpose, we assumed a  $^{12}\text{C}/^{13}\text{C}$  abundance ratio of 70 following Boogert et al. (2002) and the same rotational temperature for the two isotopologues. The population modelling of high-energy  $^{12}\text{CH}_3\text{OH}$  ( $400 < E_{up} < 1100$  K) and low-energy  $^{13}\text{CH}_3\text{OH}$  optically thin transitions allowed us to constrain the rotational temperature through the determination of the slope of the level populations, leaving only a degeneracy between  $N_{\text{tot}}$  and  $\theta_s$ . Since the optical depth  $\tau$  of each level depends on the total column density  $N_{\text{tot}}$ , low-energy optically thick transitions from  $^{12}\text{CH}_3\text{OH}$  can be used to constrain  $N_{\text{tot}}$  and  $\theta_s$ .

Table 5 presents the results of the PD analysis of the methanol population distribution. The methanol emission is relatively well modelled by the PD model for the two sources. The PD analysis converges towards one single set of input physical parameters ( $N_{\text{tot}}$ ,  $\theta_s$ ,  $T_{\text{rot}}$ ) with  $\chi^2_{\text{red}}$  of about 1.5-2 and with uncertainties up to 50 % at a one sigma level for the column densities. The PD model was able to reproduce the population of most transitions within their uncertainties except the population of some low-energy transitions of  $^{13}\text{CH}_3\text{OH}$  in IRAS2A and of  $^{12}\text{CH}_3\text{OH}$  in IRAS4A which tend to be underestimated. The rotational temperatures of methanol deduced from the PD analysis are similar in IRAS2A and IRAS4A ( $\sim 140$  K). However, the source size derived for IRAS2A is larger than for IRAS4A ( $\sim 0.36''$  versus  $\sim 0.20''$ ). The source size deduced for IRAS2A is in good agreement with the size estimated by Maret et al. (2004) but we found a smaller source size for IRAS4A, by a factor of 2.5. The source size of IRAS2A of  $0.36''$ , corresponding to a radius of 42 AU at 235 pc, is also consistent with the FWHM emission size of the  $\text{CH}_3\text{OH}$  transitions deduced by Maury et al. (2014) from elliptical gaussian fits that range between 25 AU and 70 AU.

For all other COMs, the low number of observed transitions and the relative high uncertainty on the derived column density of each level generates a degeneracy between

Table 5: Results of the PD analysis of the methanol emission.

	IRAS2A	IRAS4A
$\chi_{\text{red}}^2$	1.4	2.1
$N(\text{CH}_3\text{OH}) \text{ (cm}^{-2}\text{)}$	$5.0_{-1.8}^{+2.9} \times 10^{18}$	$1.6_{-0.8}^{+0.6} \times 10^{19}$
$N(^{13}\text{CH}_3\text{OH}) \text{ (cm}^{-2}\text{)}$	$7.1_{-2.6}^{+4.2} \times 10^{16}$	$2.3_{-1.1}^{+1.3} \times 10^{17}$
$\theta_s \text{ (")}$	$0.36_{-0.04}^{+0.04}$	$0.20_{-0.04}^{+0.08}$
$T_{\text{rot}} \text{ (K)}$	$140_{-20}^{+20}$	$140_{-30}^{+30}$

the input parameters, and prevents the PD model to converge towards one single set of input parameters: the observations are overfitted and can be reproduced by a large range of parameters giving  $\chi^2$  lower than 1. Therefore, we decided to fix the source size for the COMs emission to the size of the methanol emission, assuming that all COMs will evaporate with methanol in spite of their slightly different binding energies. Even by fixing the source size, the analysis of the glycolaldehyde and ethyl cyanide populations towards the two sources and of the ethanol population towards IRAS4A did not allow us to converge towards one set of  $N_{\text{tot}}$ , and  $T_{\text{rot}}$ . For other molecules, we were able to deduce a unique column density with relatively small uncertainties. As seen in Fig. 9 and 10, the observed population distribution is well reproduced by the best fit model since most of the column densities of upper energy levels predicted by the best fit model lie within the range of uncertainties of the observed data. The best fit models of the PDs generally consist in lower rotational temperatures and higher total column densities than the values derived with the RDs in order to reproduce the population of the optically thick transitions. For instance, the spread of the population distribution of the low upper energy levels ( $E_{\text{up}} \leq 120 \text{ K}$ ) of  $\text{CH}_3\text{OH}$ ,  $^{13}\text{CH}_3\text{OH}$ ,  $\text{HCOOCH}_3$ , or  $\text{CH}_3\text{CN}$  are explained by opacity effects. For these species, transitions showing a decrease of their population are optically thick with  $\tau$  higher than 1.



For species where the PD analysis was not able to converge towards one set of input parameters (namely glycolaldehyde, ethanol, and ethyl cyanide) and for molecules showing only one detected transition, we fixed the source size and the rotational temperature to the values found for methanol. Results of this analysis are shown in Table 6.

Most transitions of COMs whose collision rates have been computed ( $\text{CH}_3\text{OH}$ ,  $\text{CH}_3\text{CN}$ , and  $\text{HC}_3\text{N}$ ) have critical densities that range between  $10^5$  and  $10^7 \text{ cm}^{-3}$  at 100 K. They are therefore likely lower than the densities found in the hot corinos of IRAS2A and IRAS4A ( $n_{\text{H}_2} > 1.3 \times 10^8 \text{ cm}^{-3}$  following the density profiles of the two envelopes derived by Maret et al. 2004). Given the good fit to the observational data with our LTE PD analysis, it is likely that the observed species are at LTE. Most of the scattering of the population distribution can therefore be attributed to opacity effects only.

## 5. Discussion

### 5.1. Abundances in IRAS2A and IRAS4A

Methanol is likely the most abundant complex organics, and is believed to be the precursor molecule of several COMs. It is therefore worth comparing the abundance of COMs with respect to methanol to quantify the efficiency of their formation. Moreover, column densities of COMs and methanol have been derived with similar methods and from the same observational data, the estimates of the abundance ratios are therefore more accurate than the absolute abundances derived with respect to  $\text{H}_2$ . Tables 4 and 6 list the abundance ratios of the COMs with respect to methanol for the RD and PD analyses. The two targeted sources seem to have a similar chemical composition since the COMs abundance ratios differ by only a factor of two at maximum. Methyl formate is the most abundant COM of our sample, with an abundance of 1.5-3 % followed by ethanol (1-1.5 %)

Table 6: Results from the population diagram analysis for the COMs emission.

Molecule	$N_{hc}$ ( $\text{cm}^{-2}$ )	$T_{\text{rot}}$ (K)	Source size ( $''$ )	$X^a$	$X_{\text{meth}}^b$
IRAS2A					
CH <sub>3</sub> OH	$5.0^{+2.9}_{-1.8} \times 10^{18}$	$140^{+20}_{-20}$	$0.36^{+0.04}_{-0.04}$	$1.0^{+0.6}_{-0.4} \times 10^{-6}$	
<sup>13</sup> CH <sub>3</sub> OH	$7.1^{+4.2}_{-2.6} \times 10^{16}$	$140^{+20}_{-20}$	$0.36^{+0.04}_{-0.04}$	$1.4^{+0.8}_{-0.5} \times 10^{-8}$	
HCOOCH <sub>3</sub>	$7.9^{+4.6}_{-1.6} \times 10^{16}$	$160^{+50}_{-30}$	0.36 <sup>c</sup>	$1.6^{+0.9}_{-0.3} \times 10^{-8}$	$1.6^{+1.1}_{-1.0}$
CH <sub>3</sub> CN	$2.0^{+1.2}_{-0.4} \times 10^{16}$	$130^{+230}_{-40}$	0.36 <sup>c</sup>	$4.0^{+2.4}_{-0.8} \times 10^{-9}$	$0.40^{+0.28}_{-0.25}$
CH <sub>3</sub> OCH <sub>3</sub>	$5.0^{+2.9}_{-1.0} \times 10^{16}$	$110^{+60}_{-20}$	0.36 <sup>c</sup>	$1.0^{+0.6}_{-0.2} \times 10^{-8}$	$1.0^{+0.7}_{-0.6}$
C <sub>2</sub> H <sub>5</sub> OH	$7.9^{+4.6}_{-4.0} \times 10^{16}$	$270^{+230}_{-80}$	0.36 <sup>c</sup>	$1.6^{+0.9}_{-0.8} \times 10^{-8}$	$1.6^{+1.1}_{-1.2}$
HCOCH <sub>2</sub> OH	$6.8 \times 10^{15}$	$140^d$	0.36 <sup>d</sup>	$1.4 \times 10^{-9}$	$0.14^{+0.05}_{-0.08}$
C <sub>2</sub> H <sub>5</sub> CN	$1.5 \times 10^{15}$	$140^d$	0.36 <sup>d</sup>	$3.0 \times 10^{-10}$	$0.030^{+0.011}_{-0.017}$
HC <sub>3</sub> N	$9.3 \times 10^{14}$	$140^d$	0.36 <sup>d</sup>	$1.9 \times 10^{-10}$	$0.019^{+0.007}_{-0.011}$
H <sub>2</sub> <sup>13</sup> CO	$6.3 \times 10^{15}$	$140^d$	0.36 <sup>d</sup>	$1.3 \times 10^{-9}$	$0.13^{+0.05}_{-0.07}$
NH <sub>2</sub> CHO	$1.2 \times 10^{16}$	$140^d$	0.36 <sup>d</sup>	$2.4 \times 10^{-9}$	$0.24^{+0.09}_{-0.14}$
CH <sub>2</sub> CO	$6.8 \times 10^{15}$	$140^d$	0.36 <sup>d</sup>	$1.4 \times 10^{-9}$	$0.14^{+0.05}_{-0.08}$
IRAS4A					
CH <sub>3</sub> OH	$1.6^{+0.6}_{-0.8} \times 10^{19}$	$140^{+30}_{-30}$	$0.20^{+0.08}_{-0.04}$	$4.3^{+2.5}_{-2.1} \times 10^{-7}$	
<sup>13</sup> CH <sub>3</sub> OH	$2.3^{+1.3}_{-1.1} \times 10^{17}$	$140^{+30}_{-30}$	$0.20^{+0.08}_{-0.04}$	$6.2^{+3.5}_{-3.0} \times 10^{-9}$	
HCOOCH <sub>3</sub>	$5.0^{+5.0}_{-1.8} \times 10^{17}$	$60^{+20}_{-10}$	0.20 <sup>c</sup>	$1.4^{+1.4}_{-0.5} \times 10^{-8}$	$3.1^{+3.5}_{-2.1}$
CH <sub>3</sub> CN	$6.3^{+3.6}_{-1.3} \times 10^{16}$	$200^{+110}_{-40}$	0.20 <sup>c</sup>	$1.7^{+1.0}_{-0.4} \times 10^{-9}$	$0.39^{+0.30}_{-0.24}$
CH <sub>3</sub> OCH <sub>3</sub>	$1.6^{+0.9}_{-0.3} \times 10^{17}$	$80^{+40}_{-20}$	0.20 <sup>c</sup>	$4.3^{+2.5}_{-0.9} \times 10^{-9}$	$1.0^{+0.8}_{-0.6}$
C <sub>2</sub> H <sub>5</sub> OH	$1.6 \times 10^{17}$	$140^d$	0.20 <sup>d</sup>	$4.3 \times 10^{-9}$	$1.0^{+0.5}_{-0.6}$
HCOCH <sub>2</sub> OH	$4.8 \times 10^{16}$	$140^d$	0.20 <sup>d</sup>	$1.3 \times 10^{-9}$	$0.30^{+0.15}_{-0.17}$
C <sub>2</sub> H <sub>5</sub> CN	$6.4 \times 10^{15}$	$140^d$	0.20 <sup>d</sup>	$1.7 \times 10^{-10}$	$0.040^{+0.020}_{-0.023}$
HC <sub>3</sub> N	$2.9 \times 10^{15}$	$140^d$	0.20 <sup>d</sup>	$7.8 \times 10^{-11}$	$0.018^{+0.009}_{-0.010}$
H <sub>2</sub> <sup>13</sup> CO	$1.1 \times 10^{16}$	$140^d$	0.20 <sup>d</sup>	$3.0 \times 10^{-10}$	$0.069^{+0.034}_{-0.040}$
NH <sub>2</sub> CHO	$1.9 \times 10^{16}$	$140^d$	0.20 <sup>d</sup>	$5.1 \times 10^{-10}$	$0.12^{+0.06}_{-0.07}$
CH <sub>2</sub> CO	$3.4 \times 10^{16}$	$140^d$	0.20 <sup>d</sup>	$9.2 \times 10^{-10}$	$0.21^{+0.10}_{-0.12}$

Note. — a: The abundances relative to H<sub>2</sub> are obtained from  $N(\text{H}_2)$  derived at 145 GHz in Table 1 assuming an homogeneous H<sub>2</sub> column density within the beam. b: The abundances relative to CH<sub>3</sub>OH were computed from  $N(\text{CH}_3\text{OH})$  derived from the PD analysis of the CH<sub>3</sub>OH and <sup>13</sup>CH<sub>3</sub>OH emissions and adapted for the same source size. c: The source size was assumed to be equal to that of methanol when the size could not be constrained. d: The source size and rotational temperatures were assumed to be equal to

and di-methyl ether (1 %). Other COMs are detected with abundances lower than 1 %: glycolaldehyde and methyl cyanide show abundances of 0.15 and 0.40 % respectively whilst ethyl cyanide is detected with an abundance of 0.03-0.04 %.

Table 4 compares the column densities and abundance ratios deduced from the RD analysis with previous single-dish studies by Bottinelli et al. (2004a); Maret et al. (2005); Bottinelli et al. (2007) carried out towards IRAS2A and IRAS4A. The column densities obtained from these previous observations suffer from several limitations: most detected transitions have low upper energy levels and the large beam of single-dish telescopes encompasses the cold envelope where weakly excited lines may have contaminated the hot corino emission (see the interferometric maps in Figures 6 and 7). Consequently, the rotational temperatures and column densities derived with the PdBi are higher ( $T_{\text{rot}} = 80\text{-}290$  K and  $300\text{-}360$  K in this work towards IRAS2A and IRAS4A versus 100 and 25 K respectively and higher column densities up to one order of magnitude), since the interferometric observations probe material closer to the central protostars. The abundance ratios deduced from previous single-dish studies have also higher relative uncertainties than in this work due to the different telescope calibrations since the observations of methanol and COMs have been carried out separately. The abundances relative to methanol derived from our interferometric observations therefore differ from the abundances obtained with the single-dish observations, the latter being usually overestimated. The methyl formate abundance derived in IRAS2A of 2 % is consistent with the upper limit of 85 % by Bottinelli et al. (2007). However, the abundance derived in IRAS4A of  $\sim 3$  % is 18 times lower than the value derived by Bottinelli et al. (2004a) and using the column density of  $\text{CH}_3\text{OH}$  derived by Maret et al. (2005). The higher abundance of methyl formate with respect to methanol in IRAS4A derived from single-dish observations is explained by the lower column density of methanol derived in Maret et al. (2005) assuming optically thin emission. For the same reasons, the abundances of methyl cyanide derived in the two

sources by our PD analysis are also lower, by a factor of 3 to 6, than the abundances obtained by Bottinelli et al. (2007). Di-methyl ether, ethanol and ethyl cyanide have not been detected with single-dish telescopes towards IRAS2A and IRAS4A but their upper limits agree well with our observations.

We report here the first detection of glycolaldehyde in low-mass protostars other than IRAS 16293. Glycolaldehyde co-exists with its isomer methyl formate with a  $[\text{HCOOCH}_3]/[\text{HCOCH}_2\text{OH}]$  abundance ratio of  $12^{+7}_{-2}$  towards IRAS2A and of  $10^{+10}_{-4}$  towards IRAS4A. These abundance ratios are similar to the ratios of  $\sim 13$  found in the sources A and B of the IRAS 16293 protostellar binary system by Jørgensen et al. (2012) from high angular resolution ALMA observations. They are also consistent with the ratios derived towards SgrB2(N) ranging from 52 in the hot core (Hollis et al. 2001) to 5 found on more extended scales (Hollis et al. 2000).

## 5.2. Comparison with Other Sources

The COMs abundances derived in this work are compared with other published data of low-mass, intermediate-mass, and high-mass hot cores obtained with single dish and interferometric sub-mm telescopes in Figure 8 showing the COMs abundances as a function of the protostar luminosity. Table 17 lists the abundances of selected COMs towards the hot cores shown in Fig. 8 along with the references. For this purpose, we only selected observational studies where several transitions of methanol were detected. For most of the works, COMs and methanol were detected simultaneously and the abundance ratios were derived from the main isotopologue  $\text{CH}_3\text{OH}$ , either by assuming optically thin emission and LTE population (MacDonald et al. 1996; Ikeda et al. 2001; Beuther et al. 2007, 2009; Palau et al. 2011; Öberg et al. 2011, 2014), by neglecting optically thick lines in the RD analysis (Bisschop et al. 2007; Isokoski et al. 2013) or by taking the opacity of the lines

into account in their model (Nummelin et al. 2000; Qin et al. 2010; Crockett et al. 2014; Neill et al. 2014). Fuente et al. (2014) derived the methanol abundances from the  $^{13}\text{CH}_3\text{OH}$  isotopologue. For IRAS 16293, we combined the  $\text{CH}_3\text{OH}$  absolute abundance derived by Schöier et al. (2002) with the COMs abundances obtained by Jaber et al. (2014), both from a radiative transfer modelling, to obtain the abundance ratios. We also derived the glycolaldehyde abundance from Jørgensen et al. (2012).

Overall, the abundance ratios of COMs estimated towards IRAS2A and IRAS4A in this work tend to be lower than the abundances derived in other low-mass protostars from single-dish observations. For example, methyl formate and methyl cyanide have been detected towards 5 other low-mass sources and show abundances of 5-50% and 0.8-1.7% respectively, representing a factor of 1.5-15 and 1.5-3 higher than in our work. Since these observations suffer from the same limitations than the single-dish observations by Bottinelli et al. (2004a, 2007) presented in the previous section (low number of detected transitions, low upper energy levels of detected transitions, large beam encompassing the external envelope), the discrepancy likely comes from the differences in the observational methods and does not necessarily reflect differences in the chemistry between the sources.

Table 7 summarises the mean abundance ratios of COMs in low-mass, intermediate-mass and high-mass protostars. Along with Fig. 8, they allow us to investigate any possible correlation of the COMs abundances with the protostar luminosity. For each molecule, the data has been fitted by a linear curve depicted by the dashed curve whose slope  $a$  is shown at the top left of each panel of Fig. 8, by considering detected ratios only. It can be noticed that the abundances of the six COMs tend to slightly increase with the protostar luminosity. However, for all species but  $\text{C}_2\text{H}_5\text{CN}$ , the increase remains negligible compared to the dispersion of the abundance ratio values. We can conclude that the abundance ratio of these COMs stays relatively constant with the protostar luminosity within six orders of

magnitude. In spite of their lower luminosities, inducing lower temperatures, and smaller sizes, low-mass protostars seem to be as chemically complex as high-mass protostars.

Table 7: Averaged COMs abundances (in %) with respect to methanol.

Molecule	LMP	IMP	HMP	G13	R01
HCOOCH <sub>3</sub>	$14^{+41}_{-7}$	$1.3 \pm 0.2$	$14^{+18}_{-6}$	0.082 - 0.84	0.013 - 0.67
CH <sub>3</sub> OCH <sub>3</sub>	$12^{+28}_{-10}$	1.2	$30^{+23}_{-18}$	0.44 - 0.74	0.62 - 18
HCOCH <sub>2</sub> OH	$0.38^{+0.31}_{-0.18}$	/	3.1	0.54 - 1.3	-
C <sub>2</sub> H <sub>5</sub> OH	$1.3^{+0.3}_{-0.3}$	$3.2^{+5.9}_{-2.0}$	$6.4^{+6.6}_{-4.1}$	0.50 - 2.6	15 - 23
CH <sub>3</sub> CN	$1.1^{+0.4}_{-0.5}$	0.52	$5.1^{+7.8}_{-3.3}$	0.034 - 0.45	3.0 - 3.7
C <sub>2</sub> H <sub>5</sub> CN	$0.090^{+0.11}_{-0.055}$	/	$4.1^{+16}_{-2.8}$	0.052 - 0.79	-

Note. — LMP, IMP, and HMP stand for Low-Mass ( $L < 100L_{\odot}$ ), Intermediate-Mass ( $100L_{\odot} < L < 10^4L_{\odot}$ ), and High-Mass ( $10^4L_{\odot} < L$ ) Protostars. The two abundances in the column “G13” are the peak abundances of the “Fast” and “Slow” models of Garrod (2013). The two abundances in the column “R01” are the peak abundances of the models including ammonia at  $T = 100$  and  $T = 300$  K of Rodgers and Charnley (2001).

### 5.3. Comparison with Chemical Model Predictions

We compare the observed abundance ratios in low-mass, intermediate-mass, and high-mass protostars with the results of two astrochemical models in Table 7. The model of Garrod (2013) is a multilayer gas-grain astrochemical model in which COMs are assumed to be mostly formed at the surface of interstellar ices during the warm-up protostellar phase ( $30 \text{ K} < T < 100 \text{ K}$ ) through radical recombination induced by the UV photodissociation of the main ice components. In the model of Rodgers and Charnley (2001), COMs are only formed by warm gas phase chemistry for a set of constant physical parameters representative

of hot cores ( $n_{\text{H}} = 10^7 \text{ cm}^{-3}$ ;  $T = 100$  and  $300 \text{ K}$ ) after the sublimation of interstellar ices with typical ice composition, already containing  $\text{C}_2\text{H}_5\text{OH}$  with an abundance of 20 % with respect to  $\text{CH}_3\text{OH}$ . Although the rates of some key reactions for the formation of COMs have been lowered meanwhile, such as the methyl cation transfer reaction between  $\text{H}_2\text{CO}$  and  $\text{CH}_3\text{OH}_2^+$  or the electronic recombination of protonated COMs ions, this model still provides a good basis to estimate the formation efficiency of methyl formate, di-methyl ether and methyl cyanide in the gas phase.

Both models tend to underpredict the abundance of methyl formate relative to methanol observed in low-mass and high-mass protostars by at least 1-2 orders of magnitude. Moreover, the model of Garrod (2013) also underpredicts the  $[\text{HCOOCH}_3]/[\text{HCOCH}_2\text{OH}]$  abundance ratio by two orders of magnitude since it seems to reproduce well the observed abundance of glycolaldehyde. This comparison suggests that the chemical network forming methyl formate either in the gas phase or on ices is still incomplete. Possible alternative branching ratios for the photodissociation of  $\text{CH}_3\text{OH}$  on ices or other gas phase reactions involving  $\text{HCOOH}$  could enhance its formation. The observed abundances of di-methyl ether and methyl cyanide with respect to methanol can be reproduced by the gas phase model of Rodgers and Charnley (2001) only. The absolute abundances of these two molecules are similar in the two models ( $10^{-8} - 10^{-7}$  for di-methyl ether and  $10^{-9} - 10^{-8}$  for methyl cyanide), showing that warm gas phase chemistry tends to be as efficient as surface chemistry to produce these COMs. However, the difference in the abundances comes from the efficient destruction of methanol in the warm gas in the model of Rodgers and Charnley (2001) increasing the abundance ratio of COMs. No efficient formation routes in the gas phase have been proposed for glycolaldehyde and ethanol but their formation at the surface of interstellar ices seem to be efficient enough to reproduce the observations towards low-mass and high-mass protostars. The abundance of ethyl cyanide shows an increase of almost two orders of magnitude between low-mass and high-mass protostars. Grain

surface chemistry is able to reproduce the observations towards low-mass protostars but not towards high-mass hot cores. It is also possible that models also missed gas phase reactions, as it was the case for methyl formate, where a new gas phase reaction has been recently recognised by Balucani et al. (2015).

## 6. Conclusions

In this work, we have presented interferometric multi-line observations of methanol ( $\text{CH}_3\text{OH}$ ,  $^{13}\text{CH}_3\text{OH}$ ) and various COMs towards the two bright low-mass protostars NGC1333-IRAS2A and -IRAS4A, carried out with the Plateau de Bure interferometer at an angular resolution of  $\sim 2''$ . We summarize here the main conclusions of this work:

- Except for methanol, none of the transitions from complex organics are spatially resolved with the interferometer, confirming that most of the emission comes from the inner arcsec from the central protostars.
- Low-energy transitions ( $E_{\text{up}} \leq 120$  K) of methanol show extended emission around IRAS4A and trace the outflows driven by the central protostar.
- We detected a high number of transitions (up to 45 for methanol) from COMs with a wide range of upper energy levels (up to 1000 K for methanol) allowing us to accurately derive their column densities with the use of LTE methods.
- The abundances derived in this work, of a few percents for methyl formate and di-methyl ether and of  $\sim 0.5\%$  for methyl cyanide for instance, seem to be slightly lower than the abundances found towards other low-mass protostars. However, the difference likely comes from differences in the observational methods rather than different chemistries taking place in these protostars.
- We report the first detection of glycolaldehyde in low-mass protostars other than IRAS 16293, with a methyl formate-to-glycol aldehyde abundance ratio of 12 and 10 in IRAS2A



and IRAS4A, respectively.

- The analysis of the variation of the COMs abundance ratios with the protostellar luminosity shows that low-mass hot corinos seem to be as chemically rich as high-mass hot cores, in spite of their lower temperatures and their smaller sizes.
- Comparison with theoretical models shows that the two theories of COMs formation, either in the gas phase or at the surface of interstellar ices, tend to underpredict the formation of methyl formate by one to orders of magnitude. The comparison of the abundance ratios of other molecules favours a gas phase formation for di-methyl ether and methyl cyanide and a formation on ices for ethanol, ethyl cyanide, and glycolaldehyde.

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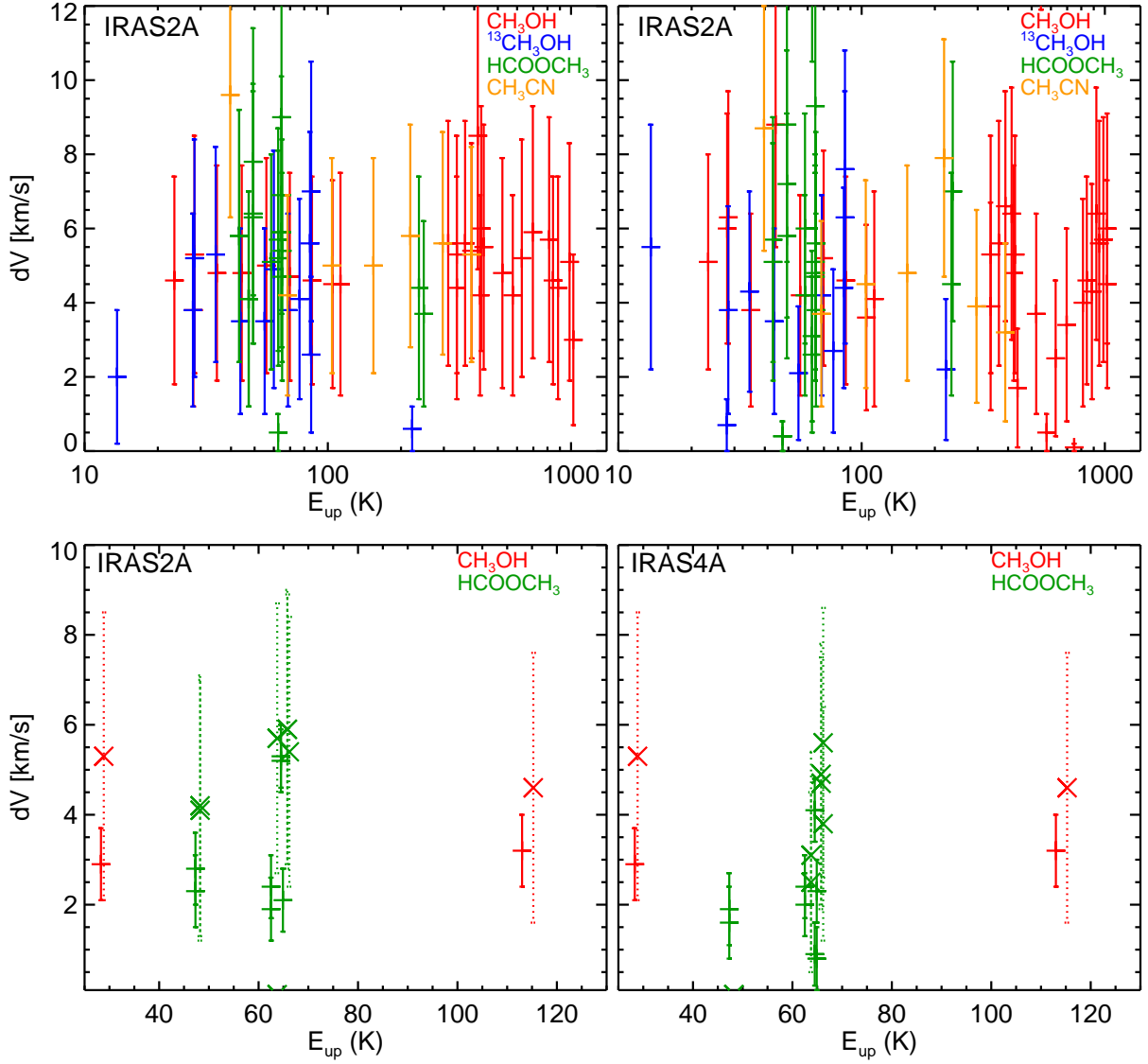


Fig. 5.— Top) Deconvolved FWHM linewidths of  $\text{CH}_3\text{OH}$ ,  $^{13}\text{CH}_3\text{OH}$ ,  $\text{HCOOCH}_3$ , and  $\text{CH}_3\text{CN}$ , obtained from gaussian fits of the WideX spectra at the source positions, as a function of  $E_{up}$ . Bottom) Comparison of the FWHM linewidths deduced from the high resolution spectra (plus signs with solid error bars) with the deconvolved FWHM linewidths from the WideX spectra (cross signs with dotted error bars) for the lines observed with the high spectral resolution correlators.

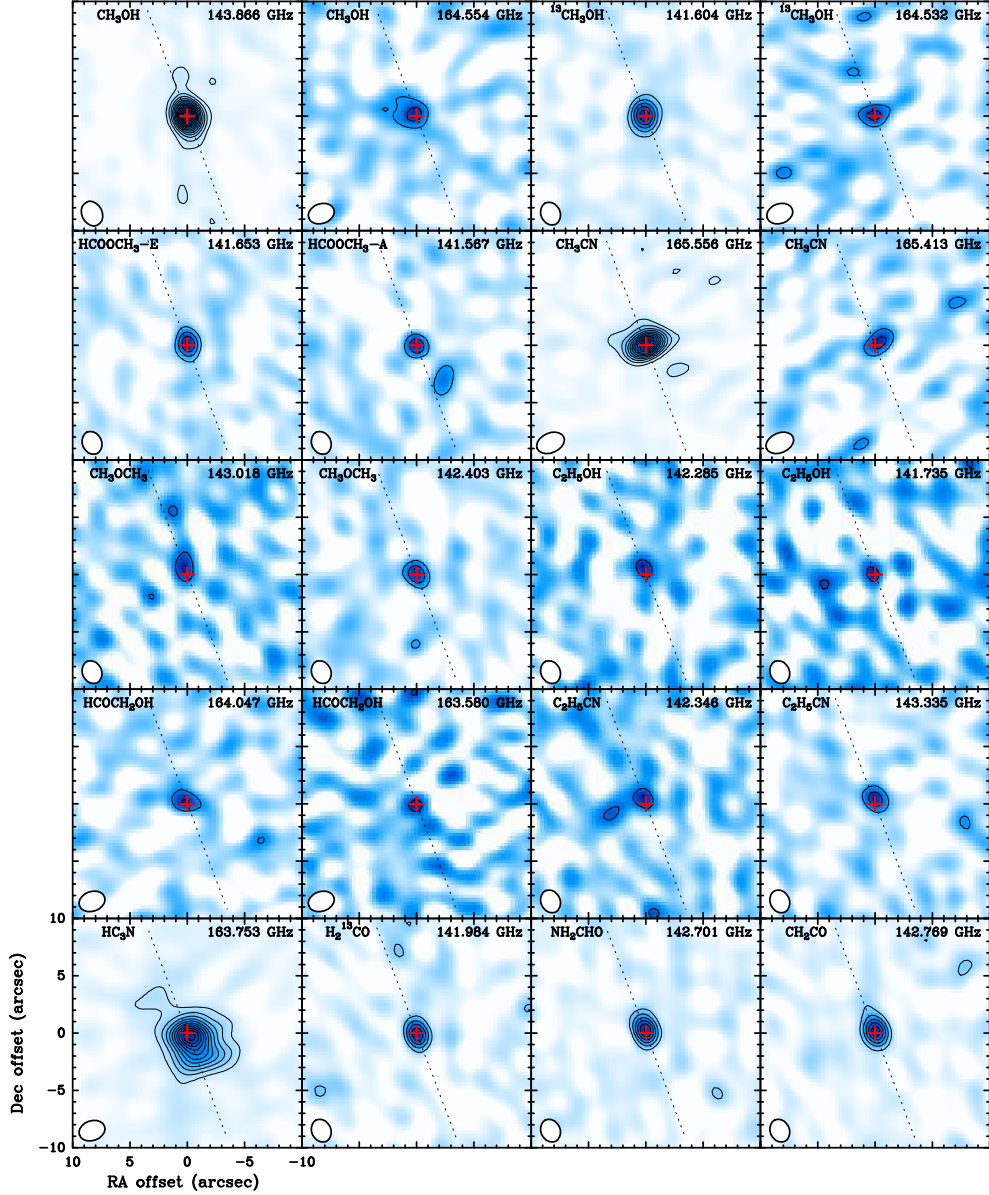


Fig. 6.— Integrated maps of selected lines of complex organic molecules towards IRAS2A. For each species where several transitions have been detected, a low-excited and a high-excited transition are shown. Contour levels are in step of  $3\sigma$ .  $1\sigma$  rms noise levels are the following: 38.5 and 18.2 mJy km/s for  $\text{CH}_3\text{OH}$ , 13.3 and 28.9 mJy km/s for  $^{13}\text{CH}_3\text{OH}$ , 13.0 and 14.6 mJy km/s for  $\text{HCOOCH}_3$ , 35.3 and 26.5 mJy km/s for  $\text{CH}_3\text{CN}$ , 15.5 and 9.64 mJy km/s for  $\text{CH}_3\text{OCH}_3$ , 11.6 and 14.5 mJy km/s for  $\text{C}_2\text{H}_5\text{OH}$ , 21.7 and 19.5 mJy km/s for  $\text{HCOCH}_2\text{OH}$ , 10.9 and 10.8 mJy km/s for  $\text{C}_2\text{H}_5\text{CN}$ , 24.8 mJy km/s for  $\text{HC}_3\text{N}$ , 13.6 mJy km/s for  $\text{H}_2^{13}\text{CO}$ , 17.6 mJy km/s for  $\text{NH}_2\text{CHO}$ , 12.0 mJy km/s for  $\text{CH}_2\text{CO}$ . The direction of the outflow is depicted by the dotted line.

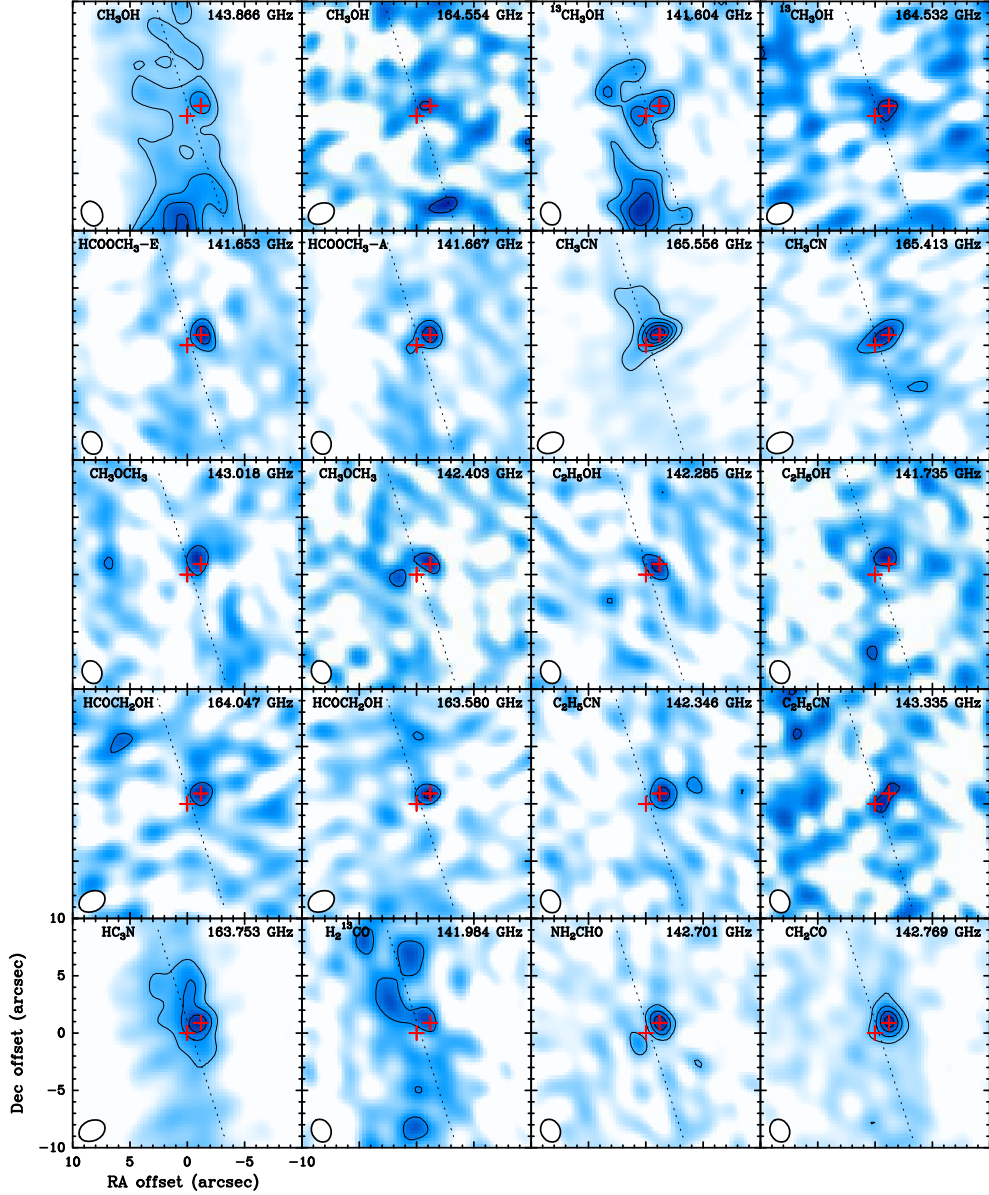


Fig. 7.— Integrated maps of selected lines of complex organic molecules towards IRAS4A. For each species where several transitions have been detected, a low-excited and a high-excited transition are shown. Contour levels are in step of  $3\sigma$ .  $1\sigma$  rms noise levels are the following: 81.9 and 22.8 mJy km/s for  $\text{CH}_3\text{OH}$ , 27.4 and 32.9 mJy km/s for  $^{13}\text{CH}_3\text{OH}$ , 25.4 and 24.5 mJy km/s for  $\text{HCOOCH}_3$ , 43.6 and 24.0 mJy km/s for  $\text{CH}_3\text{CN}$ , 16.9 and 11.4 mJy km/s for  $\text{CH}_3\text{OCH}_3$ , 12.4 and 19.8 mJy km/s for  $\text{C}_2\text{H}_5\text{OH}$ , 18.8 and 17.2 mJy km/s for  $\text{HCOCH}_2\text{OH}$ , 11.6 and 16.0 mJy km/s for  $\text{C}_2\text{H}_5\text{CN}$ , 109 mJy km/s for  $\text{HC}_3\text{N}$ , 27.8 mJy km/s for  $\text{H}_2^{13}\text{CO}$ , 14.3 mJy km/s for  $\text{NH}_2\text{CHO}$ , 22.4 mJy km/s for  $\text{CH}_2\text{CO}$ . The direction of the outflow is depicted by the dotted line.

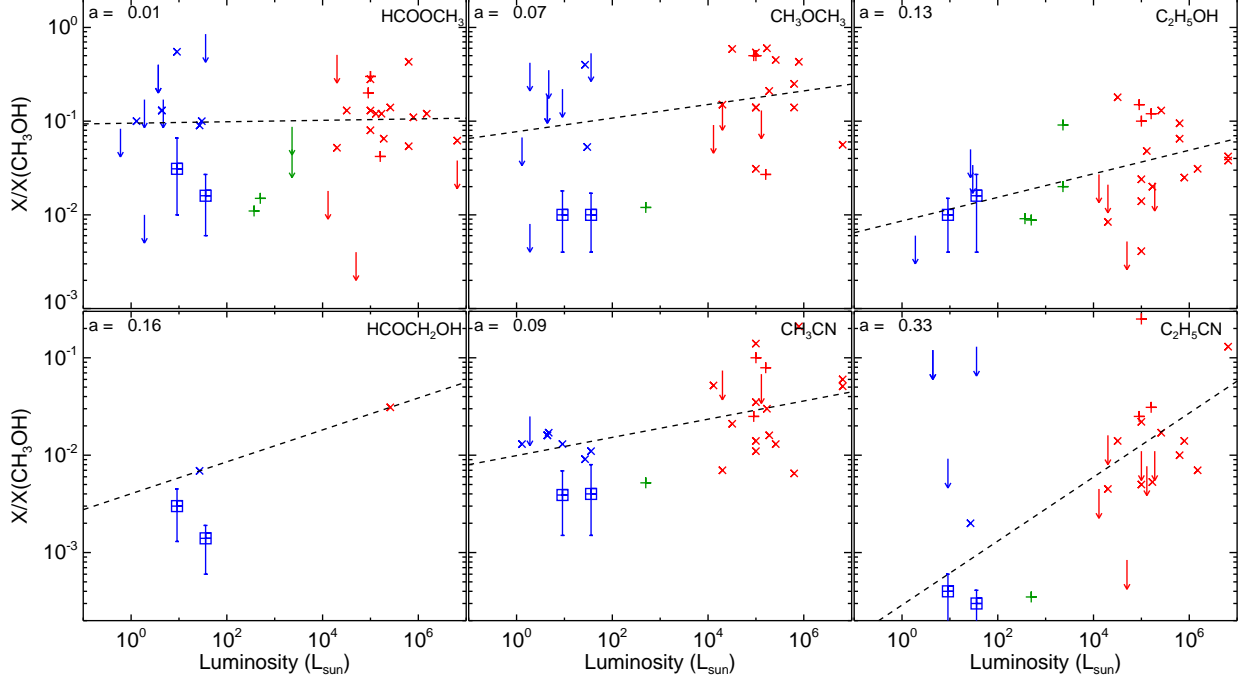


Fig. 8.— Abundances of COMs detected in several transitions with respect to  $\text{CH}_3\text{OH}$  towards a sample of low-mass ( $L < 100 L_{\odot}$ ; blue), intermediate-mass ( $100 L_{\odot} < L < 10^4 L_{\odot}$ ; green), and high-mass ( $L > 10^4 L_{\odot}$ ; red) hot cores as function of the protostar luminosity. Cross and plus symbols represent abundance ratios derived with single-dish telescopes and interferometers, respectively. The abundance ratios derived in this work with the Population Diagram analyses are denoted by squares with error bars. The sample of hot cores is listed in Table 17 along with their physical and chemical properties.

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Table 8. Line parameters of CH<sub>3</sub>OH lines observed towards IRAS2A and IRAS4A-NW.

N	Frequency (GHz)	Transition	E <sub>up</sub> (K)	A <sub>ul</sub> (s <sup>-1</sup> )	IRAS2A				IRAS4A			
					Beam size ('' × '', °)	Source size <sup>a</sup> ('' × '', °)	Flux <sup>b</sup> (Jy km/s)	dV <sub>W</sub> <sup>c</sup> (km/s)	Beam size ('' × '', °)	Source size (PA) <sup>a</sup> ('' × '', °)	Flux <sup>b</sup> (Jy km/s)	dV <sub>W</sub> <sup>c</sup> (km/s)
1	145.093754	3 <sub>0</sub> -2 <sub>0</sub> E1, v <sub>t</sub> =0	27.1	1.23(-5)	2.08×1.65 (30)	1.65×1.04 (19)	2.62±0.53	-	2.16×1.73 (25)	outflow	2.54±0.51	-
	145.097435	3 <sub>0</sub> -2 <sub>0</sub> E2, v <sub>t</sub> =0	19.5	1.10(-5)								
	145.103185	3 <sub>0</sub> -2 <sub>0</sub> A <sup>+</sup> , v <sub>t</sub> =0	13.9	1.23(-5)								
2	165.050175	1 <sub>1</sub> -1 <sub>0</sub> E1, v <sub>t</sub> =0	23.4	2.35(-5)	2.31×1.72 (110)	1.01×0.87 (11)	1.28±0.26	5.8	2.39×1.77 (114)	outflow	0.38±0.08	6.2
3	165.061130	2 <sub>1</sub> -2 <sub>0</sub> E1, v <sub>t</sub> =0	28.0	2.34(-5)	2.31×1.72 (110)	1.18×0.83 (50)	1.49±0.30	5.2	2.39×1.77 (114)	outflow	0.56±0.12	7.0
4	143.865795	3 <sub>1</sub> -2 <sub>1</sub> A <sup>+</sup> , v <sub>t</sub> =0	28.3	1.07(-5)	2.24×1.77 (25)	0.83×0.63 (10)	1.02±0.21	6.7	2.24×1.77 (25)	outflow	0.44±0.13	7.5
5	165.099240	3 <sub>1</sub> -3 <sub>0</sub> E1, v <sub>t</sub> =0	35.0	2.33(-5)	2.31×1.72 (110)	1.19×1.07 (30)	1.44±0.29	6.0	2.39×1.77 (114)	outflow	0.39±0.08	5.2
6	145.124332	3 <sub>0</sub> -2 <sub>0</sub> A <sup>-</sup> , v <sub>t</sub> =0	51.6	6.89(-6)	2.08×1.65 (30)	1.05×0.71 (32)	3.18±0.64	-	2.16×1.73 (25)	outflow	1.65±0.33	-
	145.126191	3 <sub>2</sub> -2 <sub>2</sub> E1, v <sub>t</sub> =0	36.2	6.77(-6)								
	145.126386	3 <sub>2</sub> -2 <sub>2</sub> E2, v <sub>t</sub> =0	39.8	6.86(-6)								
	145.131864	3 <sub>1</sub> -2 <sub>1</sub> E1, v <sub>t</sub> =0	35.0	1.12(-5)								
	145.133415	3 <sub>2</sub> -2 <sub>2</sub> A <sup>+</sup> , v <sub>t</sub> =0	51.6	6.89(-6)								
7	165.190475	4 <sub>1</sub> -4 <sub>0</sub> E1, v <sub>t</sub> =0	44.3	2.32(-5)	2.31×1.72 (110)	1.26×1.02 (76)	1.96±0.40	6.0	2.39×1.77 (114)	outflow	0.64±0.13	9.5
8	165.369341	5 <sub>1</sub> -5 <sub>0</sub> E1, v <sub>t</sub> =0	55.9	2.31(-5)	2.54×1.71 (113)	1.10×0.92 (42)	1.59±0.32	6.1	2.38×1.76 (114)	outflow	0.68±0.14	5.5
9	165.678649	6 <sub>1</sub> -6 <sub>0</sub> E1, v <sub>t</sub> =0	69.8	2.30(-5)	2.54×1.71 (113)	1.08×0.99 (54)	1.52±0.31	5.9	2.38×1.76 (114)	outflow	0.55±0.11	6.3
10	166.169098	7 <sub>1</sub> -7 <sub>0</sub> E1, v <sub>t</sub> =0	86.1	2.28(-5)	2.54×1.71 (113)	1.06×0.99 (30)	1.46±0.29	5.8	2.38×1.76 (114)	outflow	0.54±0.11	5.8
11	166.898566	8 <sub>1</sub> -8 <sub>0</sub> E1, v <sub>t</sub> =0	104.6	2.28(-5)	2.54×1.71 (113)	1.02×0.56 (44)	1.45±0.29	5.7	2.38×1.76 (114)	outflow	0.51±0.11	5.0
12	143.169517	7 <sub>3</sub> -8 <sub>2</sub> E1, v <sub>t</sub> =0	112.7	4.13(-6)	2.31×1.82 (26)	point	0.63±0.22	6.1	2.31×1.82 (26)	outflow	0.24±0.08	5.8
13	144.728359	3 <sub>2</sub> -2 <sub>2</sub> A <sup>+</sup> , v <sub>t</sub> =1	312.6	6.78(-6)	2.08×1.65 (30)	0.68×0.64 (-73)	0.80±0.16	6.9	2.16×1.73 (25)	point	0.23±0.05	-
	144.728782	3 <sub>2</sub> -2 <sub>2</sub> E2, v <sub>t</sub> =1	378.5	6.83(-6)								
	144.729074	3 <sub>2</sub> -2 <sub>2</sub> A <sup>-</sup> , v <sub>t</sub> =1	312.6	6.78(-6)								
14	144.733262	3 <sub>2</sub> -2 <sub>2</sub> E1, v <sub>t</sub> =1	413.8	6.80(-6)	2.08×1.65 (30)	0.38	0.96±0.19	9.4	2.16×1.73 (25)	0.48	0.42±0.09	7.6
	144.734429	3 <sub>1</sub> -2 <sub>1</sub> E1, v <sub>t</sub> =1	305.4	1.09(-5)								
	144.736349	3 <sub>0</sub> -2 <sub>0</sub> E1, v <sub>t</sub> =1	314.5	1.22(-5)								
15	144.589854	3 <sub>1</sub> -2 <sub>1</sub> A <sup>+</sup> , v <sub>t</sub> =1	339.1	1.09(-5)	2.08×1.65 (30)	0.59	0.37±0.08	6.0	2.16×1.73 (25)	0.3	0.21±0.06	6.7
16	144.878576	3 <sub>1</sub> -2 <sub>1</sub> A <sup>-</sup> , v <sub>t</sub> =1	339.2	1.09(-5)	2.08×1.65 (30)	0.35	0.43±0.09	6.7	2.16×1.73 (25)	0.91×0.75 (-5)	0.23±0.06	5.6
17	143.108385	17 <sub>0</sub> -17 <sub>-1</sub> E, v <sub>t</sub> =0	366.8	6.44(-6)	2.06×1.67 (26)	0.61×0.50 (0)	0.52±0.11	6.9	2.06×1.67 (26)	0.50	0.18±0.04	6.9
18	166.773281	5 <sub>1</sub> -5 <sub>0</sub> A <sup>+</sup> , v <sub>t</sub> =1	390.0	1.85(-5)	2.54×1.71 (113)	0.91×0.46 (67)	0.90±0.19	6.4	2.38×1.76 (114)	0.58	0.45±0.10	7.5
19	165.074355	14 <sub>6</sub> -15 <sub>5</sub> E1, v <sub>t</sub> =0	422.4	5.36(-6)	2.31×1.72 (110)	0.90×0.62 (37)	0.47±0.10	5.5	2.39×1.77 (114)	2.02×1.01 (-53)	0.30±0.06	6.0
20	144.750264	3 <sub>1</sub> -2 <sub>1</sub> E2, v <sub>t</sub> =1	427.3	1.08(-5)	2.08×1.65 (30)	point	0.28±0.06	7.2	2.16×1.73 (25)	0.15	0.12±0.04	6.7
21	144.768193	3 <sub>0</sub> -2 <sub>0</sub> A <sup>+</sup> , v <sub>t</sub> =1	437.5	1.22(-5)	2.08×1.65 (30)	0.46×0.40 (49)	0.28±0.06	6.8	2.16×1.73 (25)	1.23	0.09±0.03	4.4
22	144.572025	3 <sub>0</sub> -2 <sub>0</sub> A <sup>+</sup> , v <sub>t</sub> =2	522.1	1.42(-5)	2.08×1.65 (30)	0.52	0.21±0.04	6.3	2.16×1.73 (25)	1.61×1.03 (-60)	0.15±0.04	5.5
	144.571262	3 <sub>2</sub> -2 <sub>2</sub> E2, v <sub>t</sub> =2	658.8	6.74(-6)								
23	144.583961	3 <sub>1</sub> -2 <sub>1</sub> E2, v <sub>t</sub> =2	545.9	1.09(-5)	2.08×1.65 (30)		0.20±0.04	4.0	2.16×1.73 (25)		0.15±0.05	16.3
24	166.843662	11 <sub>2</sub> -10 <sub>3</sub> E1, v <sub>t</sub> =1	553.0	1.19(-6)	2.54×1.71 (113)	0.48	0.07±0.02	3.5	2.38×1.76 (114)	-	< 0.04	-

Table 8—Continued

25	164.299104	15 <sub>2</sub> -14 <sub>1</sub> A <sup>-</sup> , $v_t=1$	576.0	1.44(-5)	2.31×1.72 (110)	0.94×0.64 (-86)	0.43±0.10	5.5	2.39×1.77 (114)	1.95×0.94 (-83)	0.21±0.05	3.6
26	142.276432	9 <sub>7</sub> -10 <sub>6</sub> E2, $v_t=1$	627.5	1.11(-6)	2.06×1.67 (26)	1.66×0.50 (-42)	0.03±0.01	6.6	2.06×1.67 (26)	0.41	0.02±0.01	4.8
27	144.281736	3 <sub>1</sub> -2 <sub>1</sub> A <sup>+</sup> , $v_t=2$	696.6	1.07(-5)	2.08×1.65 (30)	point	0.08±0.02	7.2	2.16×1.73 (25)	0.78	0.06±0.03	5.3
28	144.530553	3 <sub>0</sub> -2 <sub>0</sub> E1, $v_t=2$	748.1	1.22(-5)	2.08×1.65 (30)	0.34	0.04±0.01	4.0	2.16×1.73 (25)	1.15	0.05±0.01	4.1
29	144.499723	3 <sub>1</sub> -2 <sub>1</sub> E1, $v_t=2$	812.8	1.05(-5)	2.08×1.65 (30)	0.96	0.03±0.01	7.0	2.16×1.73 (25)	point	0.03±0.02	5.7
30	164.486238	26 <sub>2</sub> -26 <sub>1</sub> E2, $v_t=0$	843.0	2.69(-5)	2.31×1.72 (110)	0.85×0.43 (38)	0.30±0.07	5.8	2.39×1.77 (114)	1.25	0.26±0.05	5.8
31	144.195014	24 <sub>6</sub> -25 <sub>5</sub> E2, $v_t=0$	884.4	5.16(-6)	2.08×1.65 (30)	1.76×0.69 (-59)	0.05±0.01	6.0	2.16×1.73 (25)	point	0.07±0.03	5.9
32	144.437702	10 <sub>9</sub> -10 <sub>a0</sub> E2, $v_t=1$	922.9	2.82(-6)	2.08×1.65 (30)		< 0.05	-	2.16×1.73 (25)	1.39	0.04±0.01	7.6
33	144.807264	11 <sub>9</sub> -11 <sub>a0</sub> E1, $v_t=1$	948.3	4.94(-6)	2.08×1.65 (30)		< 0.05	-	2.16×1.73 (25)	1.39×0.67 (-2)	0.07±0.02	6.9
34	145.068727	19 <sub>5</sub> -20 <sub>6</sub> E2, $v_t=1$	985.9	8.88(-6)	2.08×1.65 (30)	1.33	0.07±0.02	6.5	2.16×1.73 (25)	0.6	0.10±0.04	7.0
35	163.526070	28 <sub>4</sub> -27 <sub>5</sub> A <sup>-</sup> , $v_t=0$	1021.9	8.13(-6)	2.31×1.72 (110)	1	0.04±0.02	4.7	2.39×1.77 (114)	point	0.10±0.04	7.0
36	164.554640	28 <sub>4</sub> -27 <sub>5</sub> A <sup>+</sup> , $v_t=0$	1021.9	8.28(-6)	2.31×1.72 (110)	2.57×1.42 (28)	0.08±0.03	-	2.39×1.77 (114)	point	0.04±0.02	5.7

Note. — a: Size of the best gaussian fit to the visibilities. Gaussian fits were performed on the channel showing the peak emission. When an elliptical gaussian fit was not successful, a circular gaussian was attempted. Gaussian fits resulting in a size of 0 '' are marked by "point". Lines dominated by a molecular outflow are marked by "outflow". See text for more details.

b: Flux derived from a circular mask with a diameter equal to the major axis of the beam of the telescope given for each transition in the Table.

The errors on the fluxes were computed as the quadratic sum of the statistical error and the calibration uncertainty (about  $\sim 20$  %).

c: Non-deconvolved FWHM linewidths derived from a gaussian fit of the WideX spectra towards the source positions.

Table 9. Same as Table 8 but for  $^{13}\text{CH}_3\text{OH}$ .

N	Frequency (GHz)	Transition	$E_{up}$ (K)	$A_{ul}$ ( $s^{-1}$ )	IRAS2A				IRAS4A			
					Beam size	Source size	Flux	$dV_W$	Beam size	Source size	Flux	$dV_W$
					( $'' \times ''$ , $^\circ$ )	( $'' \times ''$ , $^\circ$ )	(Jy km/s)	(km/s)	( $'' \times ''$ , $^\circ$ )	( $'' \times ''$ , $^\circ$ )	(Jy km/s)	(km/s)
1	141.603710	$3_0-2_0$ , $A^+$ , $v_t=0$	13.6	1.15(-5)	$2.06 \times 1.67$ (26)	$1.26 \times 0.84$ (-40)	$0.142 \pm 0.032$	4.6	$2.06 \times 1.67$ (26)	$2.15 \times 1.13$ (-33)	$0.099 \pm 0.026$	6.9
2	165.575639	$2_1-2_0$ , E1, $v_t=0$	27.9	2.37(-5)	$2.54 \times 1.71$ (113)	0.48	$0.167 \pm 0.044$	5.2	$2.38 \times 1.76$ (114)	$2.27 \times 1.39$ (-62)	$0.167 \pm 0.037$	3.6
3	142.807657	$3_1-2_1$ $A^-$ , $v_t=0$	28.3	1.05(-5)	$2.06 \times 1.67$ (26)	$1.20 \times 0.78$ (18)	$0.112 \pm 0.026$	6.6	$2.06 \times 1.67$ (26)	point	$0.069 \pm 0.020$	5.6
4	165.609427	$3_1-3_0$ E1, $v_t=0$	34.6	2.36(-5)	$2.54 \times 1.71$ (113)	0.56	$0.293 \pm 0.073$	6.4	$2.38 \times 1.76$ (114)	0.56	$0.191 \pm 0.042$	5.6
5	165.690996	$4_1-4_0$ E1, $v_t=0$	43.7	2.34(-5)	$2.54 \times 1.71$ (113)	$1.04 \times 0.71$ (83)	$0.270 \pm 0.068$	5.0	$2.38 \times 1.76$ (114)	0.91	$0.172 \pm 0.040$	5.0
6	165.851224	$5_1-5_0$ E1, $v_t=0$	55.0	2.33(-5)	$2.54 \times 1.71$ (113)	0.5	$0.294 \pm 0.067$	5.0	$2.38 \times 1.76$ (114)	$2.43 \times 0.73$ (-55)	$0.165 \pm 0.036$	4.1
7	142.173740	$5_2-6_1$ E2, $v_t=0$	60	4.85(-6)	$2.06 \times 1.67$ (26)	1.25	$0.033 \pm 0.010$	6.4	-	-	-	-
8	166.128782	$6_1-6_0$ E1, $v_t=0$	68.6	2.32(-5)	$2.54 \times 1.71$ (113)	$0.99 \times 0.42$ (25)	$0.327 \pm 0.078$	5.2	$2.38 \times 1.76$ (114)	$1.74 \times 0.75$ (-88)	$0.166 \pm 0.041$	5.5
9	163.872900	$7_0-6_1$ E1, $v_t=0$	76.5	9.91(-6)	$2.31 \times 1.72$ (110)	point	$0.156 \pm 0.051$	5.4	$2.39 \times 1.77$ (114)	$1.80 \times 0.74$ (-45)	$0.151 \pm 0.037$	4.5
10	166.569486	$7_1-7_0$ E1, $v_t=0$	84.5	2.31(-5)	$2.54 \times 1.71$ (113)	$0.99 \times 0.95$ (-68)	$0.419 \pm 0.093$	6.6	$2.38 \times 1.76$ (114)	$1.60 \times 1.11$ (-60)	$0.233 \pm 0.050$	5.6
11	142.896760	$6_2-7_1$ $A^-$ , $v_t=0$	85.4	5.31(-6)	$2.06 \times 1.67$ (26)	2.24	$0.085 \pm 0.021$	8.1	$2.06 \times 1.67$ (26)	$1.45 \times 0.77$ (0)	$0.090 \pm 0.023$	7.5
12	165.280537	$6_2-7_1$ $A^+$ , $v_t=0$	85.4	7.77(-6)	$2.54 \times 1.71$ (113)	point	$0.067 \pm 0.020$	4.4	$2.38 \times 1.76$ (114)	0.76	$0.079 \pm 0.018$	8.4
13	164.531587	$13_1-12_2$ $A^-$ , $v_t=0$	222.3	9.96(-6)	$2.31 \times 1.72$ (110)	1.01	$0.121 \pm 0.047$	3.6	$2.39 \times 1.77$ (114)	2.15	$0.105 \pm 0.028$	4.2

Table 10. Same as Table 8 but for HCOOCH<sub>3</sub>.

$N$	Frequency (GHz)	Transition	$E_{up}$ (K)	$A_{ul}$ (s <sup>-1</sup> )	IRAS2A				IRAS4A			
					Beam size (''×'', °)	Source size (''×'', °)	Flux (Jy km/s)	$dV_W$ (km/s)	Beam size (''×'', °)	Source size (''×'', °)	Flux (Jy km/s)	$dV_W$ (km/s)
1	141.652995	E, 11 <sub>2,9</sub> -10 <sub>2,8</sub>	43.2	4.06(-5)	2.06×1.67 (26)	point	0.090±0.024	7.2	2.06×1.67 (26)	1.65×0.50 (-14)	0.112±0.029	7.0
2	141.667012	A, 11 <sub>2,9</sub> -10 <sub>2,8</sub>	43.2	4.06(-5)	2.06×1.67 (26)	0.2	0.069±0.018	7.2	2.06×1.67 (26)	0.93	0.119±0.027	6.6
3	143.234201	E, 12 <sub>1,11</sub> -11 <sub>1,10</sub>	47.3	4.22(-5)	2.31×1.82 (26)	point	0.073±0.028	5.8	2.31×1.82 (26)	0.44	0.091±0.035	4.1
4	143.240505	A, 12 <sub>1,11</sub> -11 <sub>1,10</sub>	47.3	4.23(-5)	2.31×1.82 (26)	point	0.072±0.023	5.8	2.31×1.82 (26)	0.52	0.111±0.036	4.1
5	142.733524	E, 13 <sub>1,13</sub> -12 <sub>1,12</sub>	49.3	4.28(-5)	2.06×1.67 (26)	0.43	0.190±0.042	7.5	2.06×1.67 (26)	1.13	0.195±0.044	7.1
	142.735139	A, 13 <sub>1,13</sub> -12 <sub>1,12</sub>	49.3	4.28(-5)								
6	142.815476	E, 13 <sub>0,13</sub> -12 <sub>0,12</sub>	49.3	4.28(-5)	2.06×1.67 (26)	0.30	0.211±0.047	7.6	2.06×1.67 (26)	1.20×0.70 (32)	0.213±0.050	8.3
	142.817021	A, 13 <sub>0,13</sub> -12 <sub>0,12</sub>	49.3	4.28(-5)								
7	142.924506	E, 13 <sub>1,13</sub> -12 <sub>0,12</sub>	49.3	6.59(-6)	2.06×1.67 (26)	point	0.045±0.012	8.8	2.06×1.67 (26)	point	0.112±0.027	9.7
	142.925911	A, 13 <sub>1,13</sub> -12 <sub>0,12</sub>	49.3	6.59(-6)								
8	164.955703	E, 13 <sub>2,11</sub> -12 <sub>2,10</sub>	58.5	6.46(-5)	2.31×1.72 (110)	0.85	0.159±0.051	6.2	2.39×1.77 (114)	1.88	0.236±0.058	5.5
9	164.968638	A, 13 <sub>2,11</sub> -12 <sub>2,10</sub>	58.5	6.46(-5)	2.31×1.72 (110)	0.4	0.124±0.039	3.5	2.39×1.77 (114)	0.38	0.209±0.047	7.0
10	163.829677	E, 14 <sub>1,13</sub> -13 <sub>1,12</sub>	62.5	6.37(-5)	2.31×1.72 (110)	1.13×0.46 (66)	0.131±0.056	6.7	2.28×1.76	0.88	0.152±0.048	4.7
11	163.835525	A, 14 <sub>1,13</sub> -13 <sub>1,12</sub>	62.5	6.37(-5)	2.31×1.72 (110)	1.82×1.36 (14)	0.129±0.039	3.6	2.28×1.76 (-70)	1.64×0.73	0.149±0.043	4.4
12	165.653657	E, 14 <sub>2,13</sub> -13 <sub>1,12</sub>	62.6	7.63(-6)	2.54×1.71 (113)	1.6	0.057±0.017	6.3	2.38×1.76 (114)	1.65	0.196±0.026	14.3
	165.657529	A, 14 <sub>2,13</sub> -13 <sub>1,12</sub>	62.6	7.63(-6)								
13	163.925845	E, 15 <sub>0,15</sub> -14 <sub>1,14</sub>	64.5	1.02(-5)	2.31×1.73	1.89×1.30 (28)	0.070±0.033	7.8	2.39×1.77 (114) (-61)	2.40×0.86	0.141±0.035	6.2
	163.927369	A, 15 <sub>0,15</sub> -14 <sub>1,14</sub>	64.5	1.02(-5)								
14	163.960387	A, 15 <sub>1,15</sub> -14 <sub>1,14</sub>	64.5	6.53(-5)	2.31×1.72 (110)	1.48×0.56 (79)	0.301±0.072	6.9	2.28×1.76	0.96	0.329±0.080	5.9
	163.961884	E, 15 <sub>1,15</sub> -14 <sub>1,14</sub>	64.5	6.53(-5)								
15	163.987455	E, 15 <sub>0,15</sub> -14 <sub>0,14</sub>	64.5	6.54(-5)	2.31×1.72 (110)	0.78×0.33 (19)	0.299±0.072	6.9	2.22×1.94	0.93	0.086±0.027	6.0
15	163.988912	A, 15 <sub>0,15</sub> -14 <sub>0,14</sub>	64.5	6.54(-5)					2.22×1.94	0.55	0.136±0.034	
16	164.022026	E, 15 <sub>1,15</sub> -14 <sub>0,14</sub>	64.5	1.03(-5)	2.31×1.72 (110)	1.81×0.76 (90)	0.118±0.046	9.7	2.39×1.77 (114)	point	0.202±0.065	10.0
	164.023416	A, 15 <sub>1,15</sub> -14 <sub>0,14</sub>	64.5	1.03(-5)								
17	164.205978	E, 13 <sub>4,9</sub> -12 <sub>4,8</sub>	64.9	5.98(-5)	2.31×1.72 (110)	1.17	0.126±0.039	5.9	2.28×1.76	0.65	0.193±0.088	6.6
18	164.223815	A, 13 <sub>4,9</sub> -12 <sub>4,8</sub>	64.9	5.98(-5)	2.31×1.72 (110)	1.29×0.90 (73)	0.138±0.041	6.5	2.28×1.76 (0)	3.02×1.74	0.202±0.074	5.2
19	142.664676	E, 12 <sub>1,11</sub> -11 <sub>1,10</sub> , $v_t=1$	234.0	4.19(-5)	2.06×1.67 (26)	-	< 0.040	-	2.06×1.67 (26)	2.65×2.24 (0)	0.037±0.009	6.1
20	142.125411	E, 13 <sub>0,13</sub> -12 <sub>0,12</sub> , $v_t=1$	236.9	4.24(-5)	2.06×1.67 (26)	1.19	0.033±0.011	6.0	2.06×1.67 (26)	-	< 0.048	-
21	142.052774	A, 13 <sub>0,13</sub> -12 <sub>0,12</sub> , $v_t=1$	236.9	4.21(-5)	2.06×1.67 (26)	-	< 0.040	-	2.06×1.67 (26)	2.24×0.56 (-29)	0.038±0.018	8.1
22	166.388878	13 <sub>3,10</sub> -12 <sub>3,9</sub> , $v_t=1$	248.0	6.51(-5)	2.54×1.71 (113)	point	0.072±0.032	5.1	2.63×1.73	-	< 0.043	-





Table 11. Same as Table 8 but for CH<sub>3</sub>CN.

$N$	Frequency (GHz)	Transition	$E_{up}$ (K)	$A_{ul}$ (s <sup>-1</sup> )	IRAS2A				IRAS4A			
					Beam size ( $'' \times ''$ , $^\circ$ )	Source size ( $'' \times ''$ , $^\circ$ )	Flux (Jy km/s)	$dV_W$ (km/s)	Beam size ( $'' \times ''$ , $^\circ$ )	Source size ( $'' \times ''$ , $^\circ$ )	Flux (Jy km/s)	$dV_W$ (km/s)
1	165.569082	9 <sub>0</sub> -8 <sub>0</sub>	39.7	2.65(-4)	2.54×1.71 (113)	0.99×0.78 (45)	1.845±0.373	10.2	2.38×1.76 (114)	1.68×1.27 (-67)	1.058±0.217	9.4
	165.565891	9 <sub>1</sub> -8 <sub>1</sub>	46.9	2.62(-4)								
2	165.556322	9 <sub>2</sub> -8 <sub>2</sub>	68.3	2.52(-4)	2.54×1.71 (113)	1.07×0.70 (-57)	0.818±0.170	5.5	2.38×1.76 (114)	0.55	0.533±0.116	5.1
3	165.540377	9 <sub>3</sub> -8 <sub>3</sub>	104.0	2.36(-4)	2.54×1.71 (113)	0.98×0.77 (65)	0.986±0.205	6.1	2.38×1.76 (114)	0.51	0.493±0.104	5.7
4	165.518064	9 <sub>4</sub> -8 <sub>4</sub>	154.0	2.13(-4)	2.54×1.71 (113)	0.87×0.79 (-26)	0.541±0.116	6.1	2.38×1.76 (114)	1.22×0.5 (-56)	0.371±0.081	6.0
5	165.489391	9 <sub>5</sub> -8 <sub>5</sub>	218.3	1.83(-4)	2.54×1.71 (113)	0.44	0.426±0.097	6.8	2.38×1.76 (114)	0.52	0.482±0.109	8.7
6	165.454370	9 <sub>6</sub> -8 <sub>6</sub>	296.8	1.47(-4)	2.54×1.71 (113)	0.74×0.45 (-80)	0.331±0.078	6.6	2.38×1.76 (114)	1.29×0.65 (-74)	0.310±0.070	5.3
7	165.413015	9 <sub>7</sub> -8 <sub>7</sub>	389.5	1.04(-4)	2.54×1.70	0.44	0.111±0.049	6.4	2.38×1.76 (114)	1.88×0.75 (-20)	0.075±0.032	4.8

Table 12. Same as Table 8 but for CH<sub>3</sub>OCH<sub>3</sub>.

N	Frequency (GHz)	Transition	E <sub>up</sub> (K)	A <sub>ul</sub> (s <sup>-1</sup> )	IRAS2A				IRAS4A			
					Beam size	Source size	Flux	dV <sub>W</sub>	Beam size	Source size	Flux	dV <sub>W</sub>
					(''×'', °)	(''×'', °)	(Jy km/s)	(km/s)	(''×'', °)	(''×'', °)	(Jy km/s)	(km/s)
1	143.017994	3 <sub>2,2</sub> -2 <sub>1,1</sub> , EA	11.1	1.09e-5	2.06×1.67 (26)	0.81	0.044±0.023	10.1	2.06×1.67 (26)	3.81×1.66 (89)	0.060±0.022	8.5
	143.018373	3 <sub>2,2</sub> -2 <sub>1,1</sub> , AE	11.1	1.09e-5								
	143.020764	3 <sub>2,2</sub> -2 <sub>1,1</sub> , EE	11.1	1.09e-5								
	143.023345	3 <sub>2,2</sub> -2 <sub>1,1</sub> , AA	11.1	1.09e-5								
2	144.856766	6 <sub>3,3</sub> -6 <sub>2,4</sub> EA	31.8	1.04(-5)	2.08×1.65 (30)	0.38	0.070±0.016	5.1	2.16×1.73 (25)	1.29×1.28 (0)	0.161±0.045	19.6
	144.858984	6 <sub>3,3</sub> -6 <sub>2,4</sub> EE	31.8	1.06(-5)								
	144.855091	6 <sub>3,3</sub> -6 <sub>2,4</sub> AE	31.8	1.06(-5)								
3	143.599420	7 <sub>3,4</sub> -7 <sub>2,5</sub> AE	38.2	1.10(-5)	2.08×1.65 (30)	1.39×0.94 (-11)	0.083±0.030	16.1	2.16×1.73 (25)	point	0.136±0.044	14.0
	143.600084	6 <sub>3,3</sub> -6 <sub>2,4</sub> EA	38.2	1.10(-5)								
	143.602993	6 <sub>3,3</sub> -6 <sub>2,4</sub> EE	38.2	1.10(-5)								
	143.606232	6 <sub>3,3</sub> -6 <sub>2,4</sub> AA	38.2	1.10(-5)								
4	141.828855	8 <sub>3,5</sub> -8 <sub>2,6</sub> , AE	45.5	1.11e-5	2.06×1.67 (26)	0.65	0.149±0.045	14.3	2.06×1.67 (26)	0.46	0.140±0.049	16.7
	141.829146	8 <sub>3,5</sub> -8 <sub>2,6</sub> , EA	45.5	1.11e-5								
	141.832261	8 <sub>3,5</sub> -8 <sub>2,6</sub> , EE	45.5	1.11e-5								
	141.835521	8 <sub>3,5</sub> -8 <sub>2,6</sub> , AA	45.5	1.11e-5								
5	143.159951	13 <sub>2,12</sub> -13 <sub>1,13</sub> , EA	88	7.75e-6	2.06×1.67 (26)	0.31	0.050±0.012	7.1	2.06×1.67 (26)	point	0.079±0.022	13.0
	143.159952	13 <sub>2,12</sub> -13 <sub>1,13</sub> , AE	88	7.75e-6								
	143.162986	13 <sub>2,12</sub> -13 <sub>1,13</sub> , EE	88	7.75e-6								
	143.166020	13 <sub>2,12</sub> -13 <sub>1,13</sub> , AA	88	7.75e-6								
6	165.208844	15 <sub>3,13</sub> -15 <sub>2,14</sub> EA	122	1.75(-5)	2.31×1.72 (110)	point	0.121±0.046	13.8	2.39×1.77 (114)	2.47×1.20 (-26)	0.115±0.025	9.5
	165.208848	15 <sub>3,13</sub> -15 <sub>2,14</sub> AE	122	1.75(-5)								
	165.211731	15 <sub>3,13</sub> -15 <sub>2,14</sub> EE	122	1.75(-5)								
	165.214617	15 <sub>3,13</sub> -15 <sub>2,14</sub> AA	122	1.75(-5)								
7	164.988708	20 <sub>3,18</sub> -19 <sub>4,15</sub> AA	204	4.60(-6)	2.31×1.72 (110)	point	0.074±0.033	5.5	2.39×1.77 (114)	-	< 0.072	-
	164.990831	20 <sub>3,18</sub> -19 <sub>4,15</sub> EE	204	4.60(-6)								
	164.992951	20 <sub>3,18</sub> -19 <sub>4,15</sub> EA	204	4.60(-6)								
	164.992958	20 <sub>3,18</sub> -19 <sub>4,15</sub> AE	204	4.60(-6)								
8	142.403201	25 <sub>1,24</sub> -25 <sub>2,23</sub> , EA	313.5	1.57e-5	2.06×1.67 (26)	point	0.051±0.013	7.2	2.06×1.67 (26)	0.41	0.035±0.011	6.0
	142.403201	25 <sub>1,24</sub> -25 <sub>2,23</sub> , AE	313.5	1.57e-5								
	142.404442	25 <sub>1,24</sub> -25 <sub>2,23</sub> , EE	313.5	1.57e-5								
	142.405682	25 <sub>1,24</sub> -25 <sub>2,23</sub> , AA	313.5	1.57e-5								



Table 13. Same as Table 8 but for C<sub>2</sub>H<sub>5</sub>OH.

$N$	Frequency (GHz)	Transition	$E_{up}$ (K)	$A_{ul}$ (s <sup>-1</sup> )	IRAS2A				IRAS4A			
					Beam size ('' $\times$ '', °)	Source size ('' $\times$ '', °)	Flux (Jy km/s)	$dV_W$ (km/s)	Beam size ('' $\times$ '', °)	Source size ('' $\times$ '', °)	Flux (Jy km/s)	$dV_W$ (km/s)
1	142.285054	9 <sub>0,9,2</sub> -8 <sub>1,8,2</sub>	37.2	1.51(-5)	2.06 $\times$ 1.67 (26)	2.74 $\times$ 0.47 (1)	0.027 $\pm$ 0.011	4.1	2.06 $\times$ 1.67 (26)	0.52	0.045 $\pm$ 0.015	8.2
2	164.900973	6 <sub>0,6,1</sub> -5 <sub>1,4,0</sub>	78.8	1.22(-5)	2.31 $\times$ 1.72 (110)	-	< 0.066	-	2.39 $\times$ 1.77 (114)	-	< 0.072	-
3	144.057496	13 <sub>3,11,2</sub> -13 <sub>2,12,2</sub>	87.9	1.80(-5)	2.08 $\times$ 1.65 (30)	point	0.031 $\pm$ 0.010	-	2.16 $\times$ 1.73 (25)	point	0.067 $\pm$ 0.018	4.1
4	141.820317	8 <sub>1,7,0</sub> -7 <sub>1,6,0</sub>	88.8	2.46e-5	2.06 $\times$ 1.67 (26)	1.16	0.034 $\pm$ 0.013	4.1	2.06 $\times$ 1.67 (26)	point	0.039 $\pm$ 0.015	6.4
5	164.626167	5 <sub>4,1,0</sub> -4 <sub>3,1,1</sub>	88.8	2.06(-5)	2.31 $\times$ 1.72 (110)	-	< 0.066	-	2.39 $\times$ 1.77 (114)	-	< 0.072	-
6	164.630894	5 <sub>4,2,0</sub> -4 <sub>3,2,1</sub>	88.8	2.06(-5)	2.31 $\times$ 1.72 (110)	-	< 0.066	-	2.39 $\times$ 1.77 (114)	-	< 0.081	-
7	144.493107	14 <sub>2,13,2</sub> -14 <sub>11,4,2</sub>	92.6	1.86(-5)	2.08 $\times$ 1.65 (30)	-	< 0.046	-	2.16 $\times$ 1.73 (25)	1.19	0.045 $\pm$ 0.021	4.1
8	166.259891	10 <sub>1,10,0</sub> -9 <sub>1,9,0</sub>	102.1	4.03(-5)	2.54 $\times$ 1.71 (113)	point	0.114 $\pm$ 0.049	7.3	2.38 $\times$ 1.76 (114)	1.08	0.118 $\pm$ 0.053	4.2
9	166.758214	10 <sub>1,10,1</sub> -9 <sub>1,9,1</sub>	106.8	4.07(-5)	2.54 $\times$ 1.71 (113)	0.95	0.043 $\pm$ 0.030	-	2.38 $\times$ 1.76 (114)	point	0.053 $\pm$ 0.020	-
10	142.083012	11 <sub>2,10,1</sub> -11 <sub>1,10,0</sub>	121	9.96(-6)	2.06 $\times$ 1.67 (26)	point	0.028 $\pm$ 0.009	16.3	2.06 $\times$ 1.67 (26)	-	< 0.040	-
11	164.511879	18 <sub>3,15,2</sub> -17 <sub>4,14,2</sub>	156.8	8.73(-6)	2.31 $\times$ 1.72 (110)	point	0.036 $\pm$ 0.021	3.6	2.39 $\times$ 1.77 (114)	-	< 0.072	-
12	164.429108	14 <sub>3,12,1</sub> -13 <sub>4,10,0</sub>	160.0	4.25(-6)	2.31 $\times$ 1.72 (110)	-	< 0.066	-	2.39 $\times$ 1.77 (114)	-	< 0.072	-
13	142.046310	20 <sub>2,18,2</sub> -20 <sub>1,19,2</sub>	185.5	1.77(-5)	2.06 $\times$ 1.67 (26)	-	< 0.040	-	2.06 $\times$ 1.67 (26)	-	< 0.040	-
13	163.515766	21 <sub>3,19,2</sub> -20 <sub>4,16,2</sub>	205.4	9.22(-6)	2.31 $\times$ 1.72 (110)	-	< 0.066	-	2.39 $\times$ 1.77 (114)	-	< 0.072	-
14	141.735490	21 <sub>4,17,2</sub> -21 <sub>3,18,2</sub>	215.5	1.72e-5	2.06 $\times$ 1.67 (26)	point	0.030 $\pm$ 0.015	8.0	2.06 $\times$ 1.67 (26)	point	0.048 $\pm$ 0.018	7.6

Table 14. Same as Table 8 but for HCOCH<sub>2</sub>OH.

$N$	Frequency (GHz)	Transition	$E_{up}$ (K)	$A_{ul}$ (s <sup>-1</sup> )	IRAS2A				IRAS4A			
					Beam size	Source size	Flux	$dV_W$	Beam size	Source size	Flux	$dV_W$
					(" × ", °)	(" × ", °)	(Jy km-s)	(km/s)	(" × ", °)	(" × ", °)	(Jy km-s)	(km/s)
1	163.951686	8 <sub>3,5</sub> -7 <sub>2,6</sub> , $v = 0$	25.6	3.68(-5)	2.31 × 1.72 (110)	-	< 0.066	-	2.39 × 1.77 (114)	-	-	-
2	143.640947	14 <sub>0,14</sub> -13 <sub>1,13</sub> , $v = 0$	53.1	8.04(-5)	2.08 × 1.65 (30)	point	0.037 ± 0.020	-	2.06 × 1.67 (26)	-	< 0.129	-
3	143.765755	14 <sub>1,14</sub> -13 <sub>0,13</sub> , $v = 0$	53.1	8.06(-5)	2.08 × 1.65 (30)	-	< 0.046	-	2.06 × 1.67 (26)	point	0.069 ± 0.035	5.6
4	164.047038	15 <sub>2,14</sub> -14 <sub>1,13</sub> , $v = 0$	66.5	9.25(-5)	2.31 × 1.72 (110)	point	0.076 ± 0.034	7.4	2.39 × 1.77 (114)	1.89	0.132 ± 0.051	4.5
5	163.542260	16 <sub>0,16</sub> -15 <sub>1,15</sub> , $v = 0$	68.3	1.21(-4)	2.31 × 1.72 (110)	point	0.037 ± 0.023	7.8	2.39 × 1.77 (114)	-	0.172 ± 0.050	8.0
6	163.580057	16 <sub>1,16</sub> -15 <sub>0,15</sub> , $v = 0$	68.3	1.21(-4)	2.31 × 1.72 (110)	point	0.038 ± 0.021	8.5	2.39 × 1.77 (114)	point	0.103 ± 0.038	5.0
7	163.697251	12 <sub>7,5</sub> -12 <sub>6,6</sub> , $v = 0$	73.1	5.60(-5)	2.31 × 1.72 (110)	-	< 0.066	-	2.39 × 1.77 (114)	point	0.071 ± 0.028	3.6
8	163.709163	12 <sub>7,6</sub> -12 <sub>6,7</sub> , $v = 0$	73.1	5.60(-5)	2.31 × 1.72 (110)	-	< 0.066	-	2.39 × 1.77 (114)	1.55	0.076 ± 0.045	3.8
9	142.784665	23 <sub>6,18</sub> -23 <sub>5,19</sub>	177	5.19(-5)	2.08 × 1.65 (30)	-	< 0.040	-	2.06 × 1.67 (26)	point	0.039 ± 0.014	5.2

Table 15. Same as Table 8 but for C<sub>2</sub>H<sub>5</sub>CN.

<i>N</i>	Frequency (GHz)	Transition	$E_{up}$ (K)	$A_{ul}$ (s <sup>-1</sup> )	IRAS2A				IRAS4A			
					Beam size ('' × '', °)	Source size ('' × '', °)	Flux (Jy km/s)	$dV_W$ (km/s)	Beam size ('' × '', °)	Source size ('' × '', °)	Flux (Jy km/s)	$dV_W$ (km/s)
1	142.346330	16 <sub>2,15</sub> -15 <sub>2,14</sub>	62.7	2.37(-4)	2.06 × 1.67 (26)	2.33	0.029 ± 0.010	12.0	2.06 × 1.67 (26)	2.51 × 1.17 (-40)	0.045 ± 0.013	6.8
2	143.529200	16 <sub>3,14</sub> -15 <sub>3,13</sub>	68.5	2.39(-4)	2.08 × 1.65 (30)	point	0.025 ± 0.009	4.1	2.16 × 1.73 (25)	-	< 0.081	-
3	144.104740	16 <sub>3,13</sub> -15 <sub>3,12</sub>	68.6	2.42(-4)	2.08 × 1.65 (30)	-	< 0.046	-	2.16 × 1.73 (25)	-	< 0.081	-
4	143.506970	16 <sub>4,13</sub> -15 <sub>4,12</sub>	76.3	2.32(-4)	2.08 × 1.65 (30)	-	< 0.046	-	2.16 × 1.73 (25)	-	< 0.081	-
5	143.535290	16 <sub>4,12</sub> -15 <sub>4,11</sub>	76.3	2.32(-4)	2.08 × 1.65 (30)	-	< 0.046	-	2.16 × 1.73 (25)	-	< 0.081	-
6	164.584755	19 <sub>0,19</sub> -18 <sub>0,18</sub>	80	3.74(-4)	2.31 × 1.72 (110)	-	< 0.066	-	2.39 × 1.77 (114)	-	< 0.072	-
8	163.948705	19 <sub>1,19</sub> -18 <sub>1,18</sub>	80.1	3.69(-4)	2.31 × 1.72 (110)	-	< 0.066	-	2.39 × 1.77 (114)	point	0.110 ± 0.044	8.5
7	143.406554	16 <sub>5,12</sub> -15 <sub>5,11</sub>	86.3	2.23(-4)	2.08 × 1.65 (30)	point	0.024 ± 0.016	5.8	2.16 × 1.73 (25)	-	< 0.081	-
	143.407188	16 <sub>5,11</sub> -15 <sub>5,10</sub>	86.3	2.23(-4)	2.08 × 1.65 (30)							-
9	143.335284	16 <sub>8,8</sub> -15 <sub>8,7</sub>	129.6	1.85(-4)	2.06 × 1.67 (26)	point	0.045 ± 0.011	7.4	2.06 × 1.67 (26)	0.17	0.037 ± 0.012	11.2
	143.335284	16 <sub>8,9</sub> -15 <sub>8,8</sub>	129.6	1.85(-4)								
	143.337710	16 <sub>7,10</sub> -15 <sub>7,9</sub>	112.9	1.99(-4)								
	143.337710	16 <sub>7,9</sub> -15 <sub>7,8</sub>	112.9	1.99(-4)								
10	143.343925	16 <sub>9,7</sub> -15 <sub>9,6</sub>	148.4	1.68(-4)	2.06 × 1.67 (26)	point	0.034 ± 0.010	7.1	2.06 × 1.67 (26)	1.43	0.045 ± 0.012	4.1
	143.343925	16 <sub>9,8</sub> -15 <sub>9,7</sub>	148.4	1.68(-4)								

Table 16. Same as Table 8 but for other molecules.

Molecule	Frequency (GHz)	Transition	$E_{up}$ (K)	$A_{ul}$ (s <sup>-1</sup> )	IRAS2A				IRAS4A			
					Beam size ('' $\times$ '', °)	Source size ('' $\times$ '', °)	Flux (Jy km/s)	$dV_W$ (km/s)	Beam size ('' $\times$ '', °)	Source size ('' $\times$ '', °)	Flux (Jy km/s)	$dV_W$ (km/s)
HC <sub>3</sub> N	163.753389	18-17	74.7	3.45(-4)	2.31 $\times$ 1.72 (110)	2.66 $\times$ 1.80 (34)	0.572 $\pm$ 0.116	6.2	2.39 $\times$ 1.77 (114)	6.03 $\times$ 1.96 (17)	0.553 $\pm$ 0.121	7.4
H <sub>2</sub> <sup>13</sup> CO	141.98374	2 <sub>0,2</sub> -1 <sub>0,1</sub>	10.2	7.25(-5)	2.06 $\times$ 1.67 (26)	0.60	0.115 $\pm$ 0.028	6.3	2.06 $\times$ 1.67 (26)	2.45 $\times$ 1.25 (0)	0.062 $\pm$ 0.022	7.3
H <sub>2</sub> C <sup>18</sup> O	143.213062	2 <sub>1,1</sub> -1 <sub>1,0</sub>	22.2	5.57(-5)	2.06 $\times$ 1.67 (26)	1.94 $\times$ 1.05 (1)	0.028 $\pm$ 0.007	-	-	-	-	-
NH <sub>2</sub> CHO	142.701479	7 <sub>7,1,7,8</sub> -6 <sub>1,6,7</sub>	30.4	2.02(-4)	2.06 $\times$ 1.67 (26)	point	0.185 $\pm$ 0.043	8.2	2.06 $\times$ 1.67 (26)	point	0.090 $\pm$ 0.025	7.4
CH <sub>2</sub> CO	142.76892	7 <sub>1,6</sub> -6 <sub>1,5</sub>	40.5	3.10(-5)	2.06 $\times$ 1.67 (26)	1.03 $\times$ 0.61 (30)	0.119 $\pm$ 0.028	5.8	2.06 $\times$ 1.67 (26)	1.80 $\times$ 1.34 (34)	0.184 $\pm$ 0.040	5.1



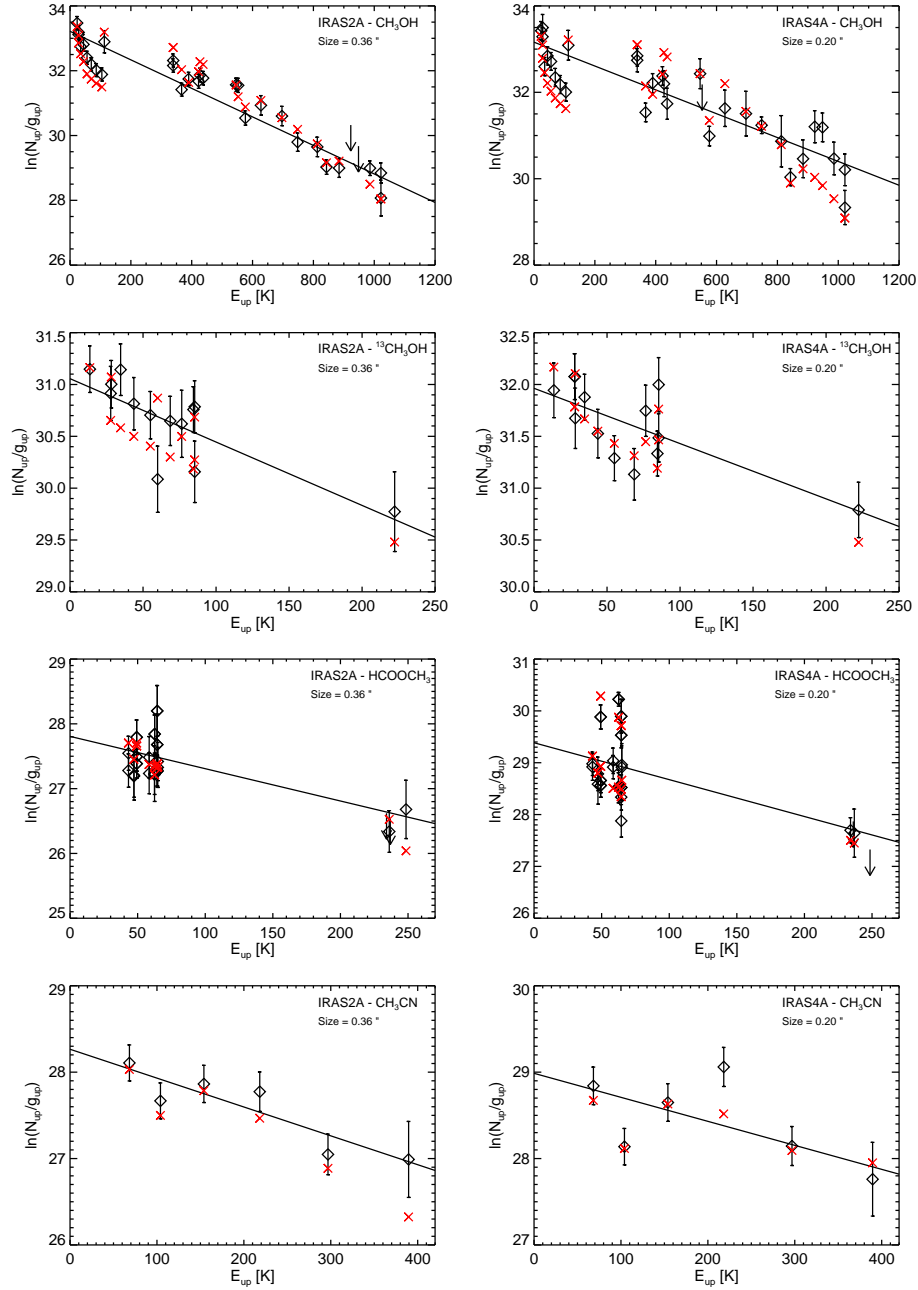


Fig. 9.— Rotational and population diagrams of methanol isotopologues ( $\text{CH}_3\text{OH}$ ,  $^{13}\text{CH}_3\text{OH}$ ),  $\text{HCOOCH}_3$ , and  $\text{CH}_3\text{CN}$  for source sizes derived from the PD analysis of the methanol population distribution ( $0.36''$  for IRAS2A and  $0.20''$  for IRAS4A). Observational data is depicted by the black diamonds. Error bars are derived assuming a calibration uncertainty of 20 % on top of the statistical error. Straight lines represent the best fit of the RD analysis to the data. Red crosses show the best fit of the PD to the data.

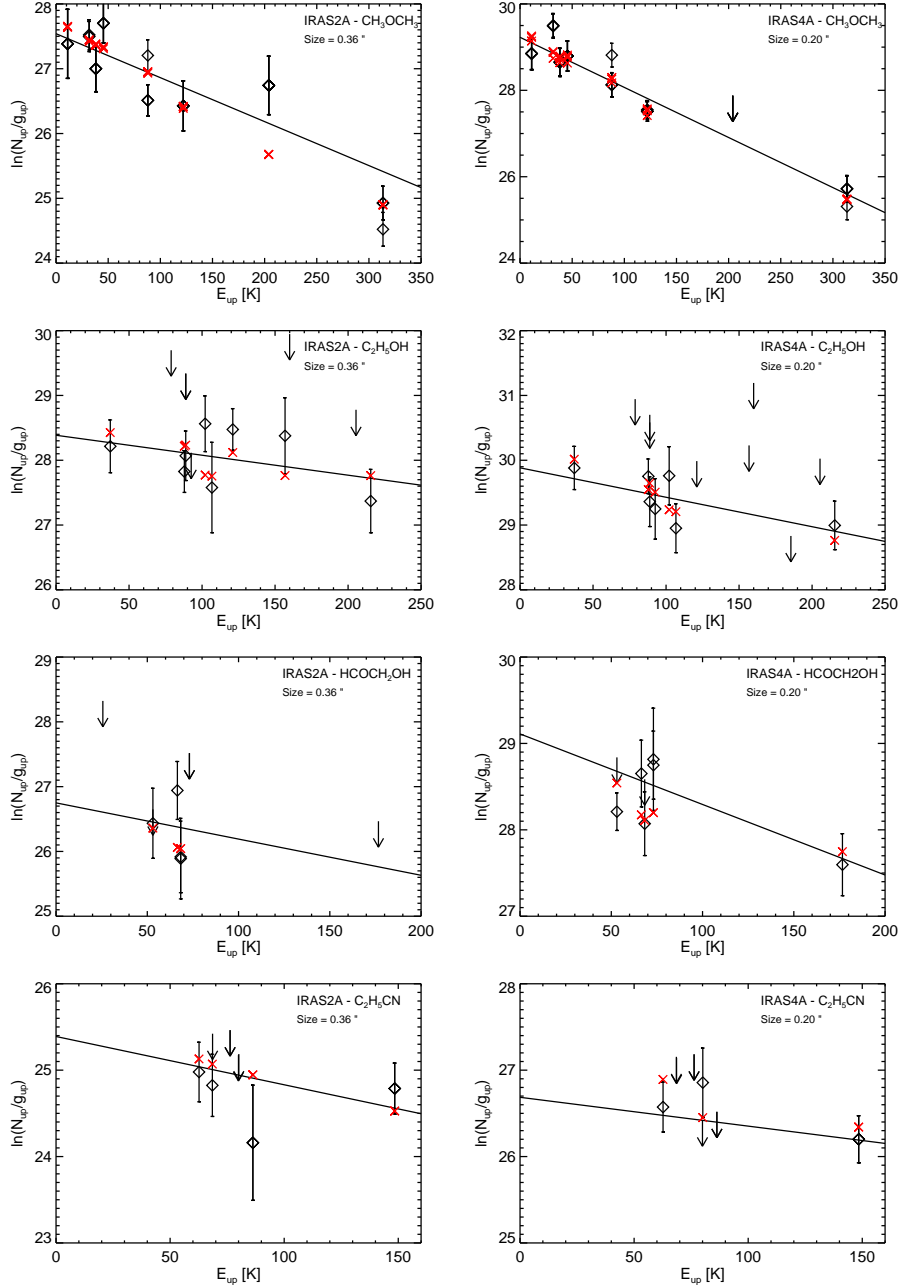


Fig. 10.— Rotational and population diagrams of  $\text{CH}_3\text{OCH}_3$ ,  $\text{C}_2\text{H}_5\text{OH}$ ,  $\text{HCOCH}_2\text{OH}$ ,  $\text{C}_2\text{H}_5\text{CN}$  for source sizes derived from the PD analysis of the methanol population distribution ( $0.36''$  for IRAS2A and  $0.20''$  for IRAS4A). Observational data is depicted by the black diamonds. Error bars are derived assuming a calibration uncertainty of 20 % on top of the statistical error. Straight lines represent the best fit of the RD analysis to the data. Red crosses show the best fit of the PD to the data.

Table 17. COMs abundances observed in high-mass, intermediate-mass, and low-mass hot  
cores.

				CH <sub>3</sub> OH	H <sub>2</sub> CO	CH <sub>2</sub> CO	HCOOCH <sub>3</sub>	HCOCH <sub>2</sub> OH	CH <sub>3</sub> OCH <sub>3</sub>	C <sub>2</sub> H <sub>5</sub> OH	CH <sub>3</sub> CN	C <sub>2</sub> H <sub>5</sub> CN	Reference
Source	$L_{\text{bol}}$ ( $L_{\odot}$ )	$N(\text{H}_2)$ ( $\text{cm}^{-2}$ )	$T_{\text{rot}}$ (K)	$X_{\text{H}_2}$	$X_{\text{meth}}$	$X_{\text{meth}}$	$X_{\text{meth}}$	$X_{\text{meth}}$	$X_{\text{meth}}$	$X_{\text{meth}}$	$X_{\text{meth}}$	$X_{\text{meth}}$	
High-mass protostars observed with single-dish telescopes													
IRAS 20126+4104	1.3(+4)	1.0(+24)	300	1.1(-7)	2.9(-1)	< 1.1(-2)	< 1.8(-2)	-	< 9.1(-2)	< 2.7(-2)	5.2(-2)	< 4.5(-3)	1
IRAS 18089-1732	3.2(+4)	1.0(+24)	300	2.0(-7)	8.6(-2)	3.1(-2)	1.3(-1)	-	5.9(-1)	1.8(-1)	2.1(-2)	1.4(-2)	1
G31.41+0.31	2.6(+5)	1.7(+23)	200	6.0(-6)	5.4(-2)	2.8(-2)	1.4(-1)	1.1(-1)	4.5(-1)	1.3(-1)	1.3(-2)	1.7(-2)	1
AFGL 2591	2.0(+4)	7.6(+22)	147	6.2(-7)	2.8(-1)	2.1(-2)	< 5.1(-1)	-	< 1.6(-1)	< 2.1(-2)	< 7.4(-2)	< 1.6(-2)	2
NGC 7538 IRS1	1.3(+5)	2.1(+23)	156	5.7(-7)	2.1(-1)	5.3(-2)	1.2(-1)	-	< 1.3(-1)	4.8(-2)	< 6.8(-2)	< 7.7(-3)	2
G24.78	7.9(+5)	4.0(+23)	211	7.0(-7)	2.3(-1)	5.2(-2)	1.1(-1)	-	4.3(-1)	2.5(-2)	2.1(-1)	1.4(-2)	2
G75.78	1.9(+5)	1.2(+23)	113	9.2(-7)	2.0(-1)	6.2(-2)	6.5(-2)	-	2.1(-1)	< 2.2(-2)	1.6(-2)	< 1.1(-2)	2
W33A	1.0(+5)	2.6(+23)	259	7.7(-7)	2.7(-1)	4.9(-2)	1.3(-1)	-	1.4(-1)	2.4(-2)	1.4(-1)	< 1.1(-2)	2
NGC 6334 IRS1	1.7(+5)	2.4(+23)	178	4.0(-6)	1.3(-1)	2.0(-2)	1.2(-1)	-	6.0(-1)	2.0(-2)	3.0(-2)	5.3(-3)	2
W3 (H <sub>2</sub> O)	2.0(+4)	1.8(+23)	181	5.6(-6)	1.8(-1)	1.5(-2)	5.2(-2)	-	1.5(-1)	8.4(-3)	7.0(-3)	4.5(-3)	2
Sgr B2 (M)	6.5(+6)	3.5(+24)	150	7.4(-9)	1.0(-1)	1.9(-2)	6.2(-2)	-	-	4.2(-2)	6.0(-2)	-	3
Sgr B2 (N)	6.5(+6)	8.0(+24)	170	6.3(-7)	4.8(-2)	1.4(-2)	< 3.8(-2)	-	5.6(-2)	3.8(-2)	5.1(-2)	1.3(-1)	4
G327.3-0.6	1.0(+5)	3.0(+24)	118	2.0(-5)	3.6(-5)	-	8.0(-2)	-	5.4(-1)	4.1(-3)	3.5(-2)	2.2(-2)	5
Orion KL - HC	1.0(+5)	3.1(+23)	128	2.2(-6)	5.5(-2)	-	-	-	3.1(-2)	-	1.4(-2)	5.0(-3)	6
Orion KL - CR	1.0(+5)	3.9(+23)	140	1.2(-6)	3.7(-2)	4.3(-3)	2.8(-1)	-	1.4(-1)	1.4(-2)	1.1(-2)	-	6
G34.3+0.15	6.3(+5)	5.3(+23)	336	7.0(-8)	-	1.8(-2)	4.3(-1)	-	2.5(-1)	9.5(-2)	6.5(-3)	-	7
G34.3+0.2	6.3(+5)	1.6(+23)	96	1.7(-7)	-	-	5.4(-2)	-	1.4(-1)	6.5(-2)	-	1.0(-2)	8
DR21(OH)	5.0(+4)	2.5(+24)	150	1.0(-8)	-	-	< 4.0(-3)	-	-	< 5.2(-3)	-	< 8.4(-4)	8
W51	1.5(+6)	3.3(+23)	208	3.0(-7)	-	-	1.2(-1)	-	-	3.1(-2)	-	7.0(-3)	8
High-mass hot cores observed with interferometers													
Orion KL	1.0(+5)	4.4(+24)	200	4.5(-9)	-	-	3.0(-1)	-	5.0(-1)	1.0(-1)	1.0(-1)	2.5(-1)	9
G29.96	9.0(+4)	3.3(+24)	200	1.2(-7)	-	-	2.0(-1)	-	5.0(-1)	1.5(-1)	2.5(-2)	2.5(-2)	10
G19.61-0.23	1.6(+5)	8.4(+23)	151	6.2(-7)	-	-	4.2(-2)	-	2.7(-2)	1.2(-1)	7.9(-2)	3.1(-2)	11
Intermediate-mass hot cores observed with interferometers													
I22198-MM2	370	2.0(+25)	120	1.2(-6)	-	-	1.1(-2)	-	-	9.1(-3)	-	-	12
A5142-MM1	2300	1.0(+25)	210	2.3(-7)	-	-	< 8.7(-2)	-	-	9.1(-2)	-	-	12
A5142-MM2	2300	2.0(+25)	140	2.0(-7)	-	-	< 5.0(-2)	-	-	2.0(-2)	-	-	12
NGC 7129 FIRS2	500	2.5(+24)	238	1.0(-6)	1.6(-2)	-	1.5(-2)	-	1.2(-2)	8.8(-3)	5.2(-3)	3.5(-4)	13
Low-mass protostars observed with single-dish telescopes													
IRAS 16293	27	2.0(+23)	84	1.0(-7)	2.6(-1)	1.0(-3)	9.0(-2)	6.9(-3)	4.0(-1)	< 5.0(-2)	9.1(-3)	< 2.0(-3)	14
IRAS2A	36	2.1(+23)	101	8.8(-7)	2.9(-1)	-	< 8.5(-1)	-	< 5.3(-1)	-	1.1(-2)	< 1.3(-1)	15
IRAS4A	9.1	1.6(+24)	24	1.4(-7)	1.4(-1)	-	5.5(-1)	-	< 2.2(-1)	-	1.3(-2)	< 9.2(-3)	15
IRAS4B	4.4	8.1(+22)	34	6.9(-6)	2.0(-1)	-	1.3(-1)	-	< 1.9(-1)	-	1.6(-2)	< 1.2(-1)	15
SMM1	30	1.3(+23)	16	1.9(-9)	-	-	1.0(-1)	-	5.3(-2)	< 3.4(-2)	-	-	16

Table 17—Continued

SMM4	1.9	1.1(+23)	13	9.5(-9)	-	-	< 1.0(-2)	-	< 8.0(-3)	< 6.0(-3)	-	-	16
B1-a	1.3	1.9(+22)	15	6.4(-10)	-	-	1.0(-1)	-	< 6.7(-2)	-	1.3(-2)	-	17
SVS 4-5	38	5.7(+22)	20	3.9(-9)	-	-	4.5(-2)	-	1.0(-1)	-	7.7(-3)	-	17
B5 IRS1	4.7	2.3(+22)	17	1.0(-9)	-	-	< 1.7(-1)	-	< 3.5(-1)	-	1.7(-2)	-	17
IRAS 03235	1.9	1.4(+23)	18	8.6(-11)	-	-	< 1.7(-1)	-	< 4.2(-1)	-	< 2.5(-2)	-	17
IRAS 04108	0.62	2.9(+22)	9	4.1(-10)	-	-	< 8.3(-2)	-	-	-	-	-	17
L1489 IRS	3.7	4.3(+22)	8	1.2(-10)	-	-	< 4.0(-1)	-	-	-	-	-	17
<b>Low-mass hot corinos observed with interferometers</b>													
IRAS2A-RD	36	2.0(+24)	179	2.5(-7)	8.1(-2)	1.4(-3)	1.9(-2)	1.5(-3)	1.2(-2)	1.5(-2)	3.0(-3)	2.7(-4)	18
IRAS2A-PD	36	2.0(+24)	140	1.0(-6)	-	-	1.6(-2)	1.4(-3)	1.0(-2)	1.6(-2)	4.0(-3)	3.0(-4)	18
IRAS4A-RD	9.1	1.4(+25)	300	1.7(-8)	6.3(-2)	2.8(-3)	1.5(-2)	2.5(-3)	8.7(-3)	1.2(-2)	1.8(-3)	4.2(-4)	18
IRAS4A-PD	9.1	1.4(+25)	140	4.3(-7)	-	-	3.1(-2)	3.0(-3)	1.0(-2)	1.0(-2)	3.9(-3)	4.0(-4)	18

Note. — 1: Isokoski et al. (2013); 2: Bisschop et al. (2007); 3: Nummelin et al. (2000); 4: Neill et al. (2014); 5: Gibb et al. (2000); 6: Crockett et al. (2014); 7: MacDonald et al. (1996); 8: Ikeda et al. (2001); 9: Beuther et al. (2009); 10: Beuther et al. (2007); 11: Qin et al. (2010); 12: Palau et al. (2011); 13: Fuente et al. (2014); 14: Maret et al. (2005); Jaber et al. (2014); 15: Maret et al. (2004, 2005); Bottinelli et al. (2004a, 2007); 16: Öberg et al. (2011); 17: Öberg et al. (2014); 18: This work